

Emi Masaki Phys.Rev.D96,043519 (2017)

with Arata Aoki, and Jiro Soda

1. Introduction

Photon - Axion Conversion

It is well known that the axions can be converted into photons and vice versa in the presence of background magnetic fields.

Cosmological Magnetic Field

$10^{-15} \text{ G} \lesssim$ Strength of Inter Galactic Magnetic Field $\lesssim 1 \text{ nG}$

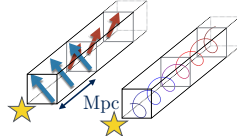
Astrophysical Origin :

vs. Ejection from Galaxy ... galactic wind, AGN jet etc.

Primordial Origin :

Inflationary Generation

... Magnetic field can have nonzero helicity.



→ We study configuration dependence of Photon-Axion Conversion.

2. Photon-Axion Conversion

$$S = \int d^4x \left[-\frac{1}{2} (\partial_\mu a \partial^\mu a + m_a^2 a^2) - \frac{1}{4} F_{\mu\nu} F^{\mu\nu} - \frac{1}{4} g_{a\gamma\gamma} a F_{\mu\nu} \tilde{F}^{\mu\nu} \right]$$

Klein-Gordon eq.

$$(\square - m_a^2) a = -g_{a\gamma\gamma} B_T \omega A_{\parallel}$$

Axion can mix with the only component **parallel** to background magnetic field.

Conversion Probability

$$P = \left(\frac{2\Delta_M}{\Delta_{osc}} \right)^2 \sin^2 \left(\frac{\Delta_{osc} z}{2} \right)$$

Axion Intensity : $I_a(z) = P \times I(z_0)$

Photon Intensity : $I(z) = [1 - P] \times I(z_0)$

$$\Delta_a = \frac{m_a^2}{2\omega} \quad \Delta_p = \frac{\omega_p^2}{2\omega} \quad \Delta_M = -\frac{1}{2} g_{a\gamma\gamma} B_T \quad \Delta_{osc} = \sqrt{(\Delta_a - \Delta_p)^2 + (2\Delta_M)^2}$$

mass term mixing term

Maxwell eq.

$$\square A_{\parallel} = -g_{a\gamma\gamma} B_T \omega a$$

$$\square A_{\perp} = 0$$

3. Astrophysical Configuration

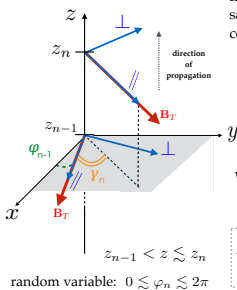
Domain Network

Within each domain, the direction of magnetic field is constant, but between domains its direction randomly changes.

$$\rho(z) \equiv \begin{pmatrix} a(z) \\ A_{\parallel}(z) \\ A_{\perp}(z) \end{pmatrix} \otimes \begin{pmatrix} a^*(z) \\ A_{\parallel}^*(z) \\ A_{\perp}^*(z) \end{pmatrix} = \frac{1}{2} \begin{pmatrix} 2I_a & K - iL & M - iN \\ I + Q & U - iV & \\ U + iV & I - Q & \end{pmatrix}$$

$$\rho(z) = \mathbf{T}(z, z_0) \rho(z_0) \mathbf{T}^\dagger(z, z_0) \quad \mathbf{T}(z_0, z_0) = 1, \quad \mathbf{T}^\dagger \mathbf{T} = 1$$

Stokes Parameters



In the limit $n = z/s \rightarrow \infty$, the conversion probability saturates so that on average one third of all photons converts to axions.

$$I_a \rightarrow \frac{1}{3}, \quad I \rightarrow \frac{2}{3}$$

Domain Numbers
 $n = z/s \rightarrow \infty$

We can find the same limit for other quadratic quantity.

$$Q = U = V = K = L = M = N = 0$$

$$I^2 = Q^2 + U^2 + V^2 \quad \Pi_L = \frac{\sqrt{Q^2 + U^2}}{I} \quad \Pi_C = \frac{|V|}{I}$$

Linear Polarization Circular Polarization

Mean-Square Values of Polarization

$\rho^2(z)$ obeys the same equation as $\rho(z)$.

$$\rho^2(z) = \mathbf{T}(z, z_0) \rho^2(z_0) \mathbf{T}^\dagger(z, z_0)$$

This implies the corresponding components of $\rho^2(z)$ have the same limit as $\rho(z)$.

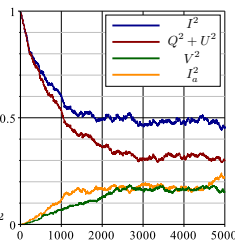
$$\bar{I}_a \equiv I_a^2 + \frac{K^2 + L^2 + M^2 + N^2}{4} = \frac{1}{3}$$

$$\bar{I} \equiv \frac{I^2 + Q^2 + U^2 + V^2}{2} + \frac{K^2 + L^2 + M^2 + N^2}{4} = \frac{2}{3}$$

Assuming equipartition for $I_a^2, Q^2, U^2, V^2, K^2, L^2, M^2, N^2$

$$\rightarrow I_a^2 = \frac{1}{6}, \quad I^2 = \frac{1}{2}, \quad Q^2 = \frac{1}{6}, \quad U^2 = \frac{1}{6}, \quad V^2 = \frac{1}{6}$$

$$\rightarrow \Pi_L = \sqrt{\frac{Q^2 + U^2}{I^2}} = 0.816 \quad \Pi_C = \sqrt{\frac{V^2}{I^2}} = 0.577$$



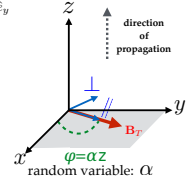
4. Helical Configuration

Magnetic Field Helicity

$$H \equiv \frac{1}{V} \int_V \mathbf{A} \cdot \mathbf{B} d^3x \quad \mathbf{B}_T = B_T \cos(\alpha z) \hat{e}_x + B_T \sin(\alpha z) \hat{e}_y$$

$$\frac{d}{dz} \begin{pmatrix} a(z) \\ A_{\parallel}(z) \\ A_{\perp}(z) \end{pmatrix} = \begin{pmatrix} \Delta_a & \Delta_M & 0 \\ \Delta_M & 0 & i\alpha \\ 0 & -i\alpha & 0 \end{pmatrix} \begin{pmatrix} a(z) \\ A_{\parallel}(z) \\ A_{\perp}(z) \end{pmatrix}$$

Wang & Lai, JCAP06(2016)006 their random variable: $0 \lesssim \varphi_n \lesssim 2\pi$

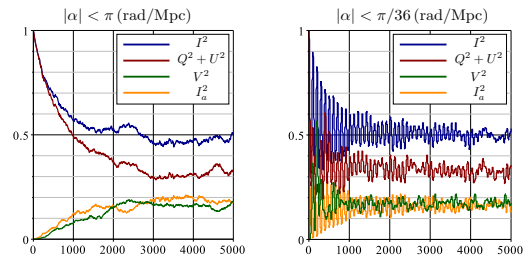


Mean-Square Values of Polarization

Left Panel : Quite similar behavior to astrophysical configuration.

Right Panel :

Convergence to the same asymptotic values via damped oscillating behavior. In case of $\alpha=0$, mean-square values oscillate characterized by conversion parameters Δ . After many domains, change of magnetic configuration is accumulated and ends up with the same asymptotic values.



5. Constraint for Physical Quantity

Observation of GRBs

The measurement of polarization says that the circular polarization of GRBs is less than 1%. Since the photon-axion conversion can produce the circular polarization, we can constrain the conversion parameters Δ so as not to contradict observation.

Mixing Term :

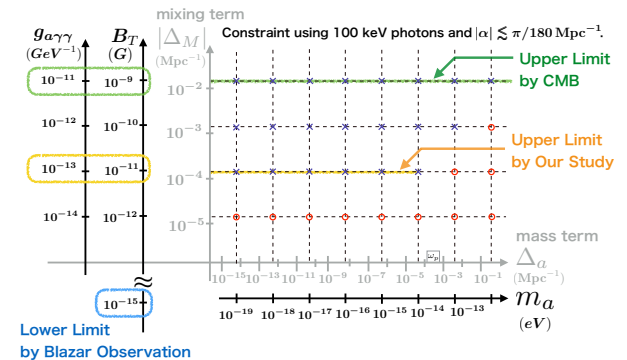
$$\Delta_M \equiv -\frac{1}{2} g_{a\gamma\gamma} B_T = -1.5 \times 10^{-2} \left(\frac{g_{a\gamma\gamma}}{10^{-11} \text{ GeV}^{-1}} \right) \left(\frac{B_T}{\text{nG}} \right) \text{ Mpc}^{-1}$$

Mass Term :

$$\Delta_a \equiv \frac{m_a^2}{2\omega} = 7.8 \times 10^{23} \left(\frac{m_a}{\text{eV}} \right)^2 \left(\frac{100 \text{ keV}}{\omega} \right) \text{ Mpc}^{-1}$$

$$\Delta_p \equiv \frac{\omega_p^2}{2\omega} = 1.1 \times 10^{-4} \left(\frac{100 \text{ keV}}{\omega} \right) \left(\frac{n_e}{10^{-7} \text{ cm}^{-3}} \right) \text{ Mpc}^{-1}, \quad \omega_p^2 \equiv e^2 \frac{n_e}{m_e}$$

For the parameters corresponding to the boundary between allowed and excluded regions, we found the mean-square values do not reach the asymptotic values. Rather, the transient regime is relevant.



6. Conclusion

- We analytically obtained asymptotic values of the variance of photon polarization in astrophysical magnetic field configuration.
- It turned out that the asymptotic behavior does not depend on the magnetic field configuration. However, when helicity is small, we showed numerically that the oscillation appears in the early phase of evolution.
- We also applied our results to the astrophysical situation. Since the constraints for physical quantity depends on magnetic field configuration, it may be possible to restrict configuration by combing with other axion or magnetic field observations.