Learning about new physics from Gravitational Waves

... and a little bit about myself

Cem Eröncel (DESY), Building 1a O1.138 14.02.2020

FH Fellow's Meeting 2020



Born in Istanbul, Turkey.



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Academic Career:

- Bsc. Electrical Engineering @ Istanbul Technical University
- Msc. Physics @ Bogazici University, Istanbul



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Hobbies:

- (Landscape) photography
- Hiking



My current work

Gravitational waves







Axions and Axion-Like-Particles (ALPs)

One of the hints of BSM: Strong CP Problem

$$\mathcal{L}_{\mathsf{SM}} \supset rac{ heta g_{\mathsf{s}}^2}{32\pi^2} G \, \widetilde{G}, \quad heta \lesssim \mathbf{10}^{-\mathbf{10}}.$$

The most popular solution is to add a new global symmetry to the SM.

 $\mathcal{G}(\mathsf{SM}) \otimes U(1)_{\mathsf{PQ}} \nrightarrow \mathcal{G}(\mathsf{SM}) \quad @f_a$

Axions are the NGBs of the broken $U(1)_{PQ}$.

$$\varphi = \underbrace{\chi}_{\langle \chi \rangle = \frac{f_a}{\sqrt{2}}} \exp \left\{ i \frac{\phi}{f_a} \right\}$$

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At low energy, non-perturbative physics (NP) introduces a periodic potential for the axion.

$$V \supset \Lambda\left[1 - \cos\left(rac{\phi}{f_{s}}
ight)
ight], \quad \Lambda \equiv {\sf NP} \; {\sf scale}$$

From this potential, axion gets a mass and couplings to other SM fields.

$$m^2 = rac{\Lambda^4}{f_a^2}$$
 , $rac{g_{\phi\gamma\gamma}}{4}rac{a}{f}F\widetilde{F} \sim g_{\phi\gamma\gamma}rac{\phi}{f_a}{f E}\cdot{f B}.$

If axions solve Strong CP, then $\Lambda \sim m_{\pi}^2 f_{\pi}^2$.

$$\Lambda \sim egin{cases} m_\pi^2 f_\pi^2, & {\sf QCD} \ {\sf Axion} \ {\sf free \ parameter}, & {\sf ALP} \end{cases}$$

ALP parameter space and the constraints: PDG(2019)



Axion Fragmentation

Efficient production of axion quanta is possible when an axion-like field ϕ rolls down a wiggly potential.

$$V(\phi) = -\mu^3 \phi + \Lambda^4 \cos\left(rac{\phi}{f}
ight)$$

Fonseca, Morgante, Sato, Servant, 1911.08472 & 1911.08473 These axion quanta source gravitational waves:

$$ar{h}_{ij}^{\prime\prime}(\mathbf{k})+\left(k^2-rac{a^{\prime\prime}}{a^\prime}
ight)ar{h}_{ij}(\mathbf{k})=rac{2a^3}{M_{
m pl}^2}\Pi_{ij}^{ au au}.$$

For this and other related models, the $\ensuremath{\mathsf{peak}}$ GW spectral density $\ensuremath{\mathsf{scales}}$ as

$$h^2 \Omega_{
m GW}(k_{
m peak}) \sim \left(rac{f_a}{M_{
m pl}}
ight)^n$$



Ongoing project with Géraldine Servant, Philip Sørensen and Ryosuke Sato.

My Favourite Plot

GW sensitivity curves and sources



gwplotter.com

Thank you!

Contact:

DESY. Deutsches Elektronen-Synchrotron www.desy.de Cem Eröncel o 0000-0002-9308-1449 Theory cem.eroncel@desy.de