HEP2023
HAMBURG

Global Network of Cavities to Search for Gravitational Waves: GravNet

A novel scheme to hunt gravitational waves signatures from the early universe

EPS-HEP, Hamburg, 2023

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Cluster of Excellence PRiSMA+ | Precision Physics, Fundamental Interactions and Structure of Matter

JG|U

Based on [arXiv:2308.11497]

Introduction - Gravitational Waves

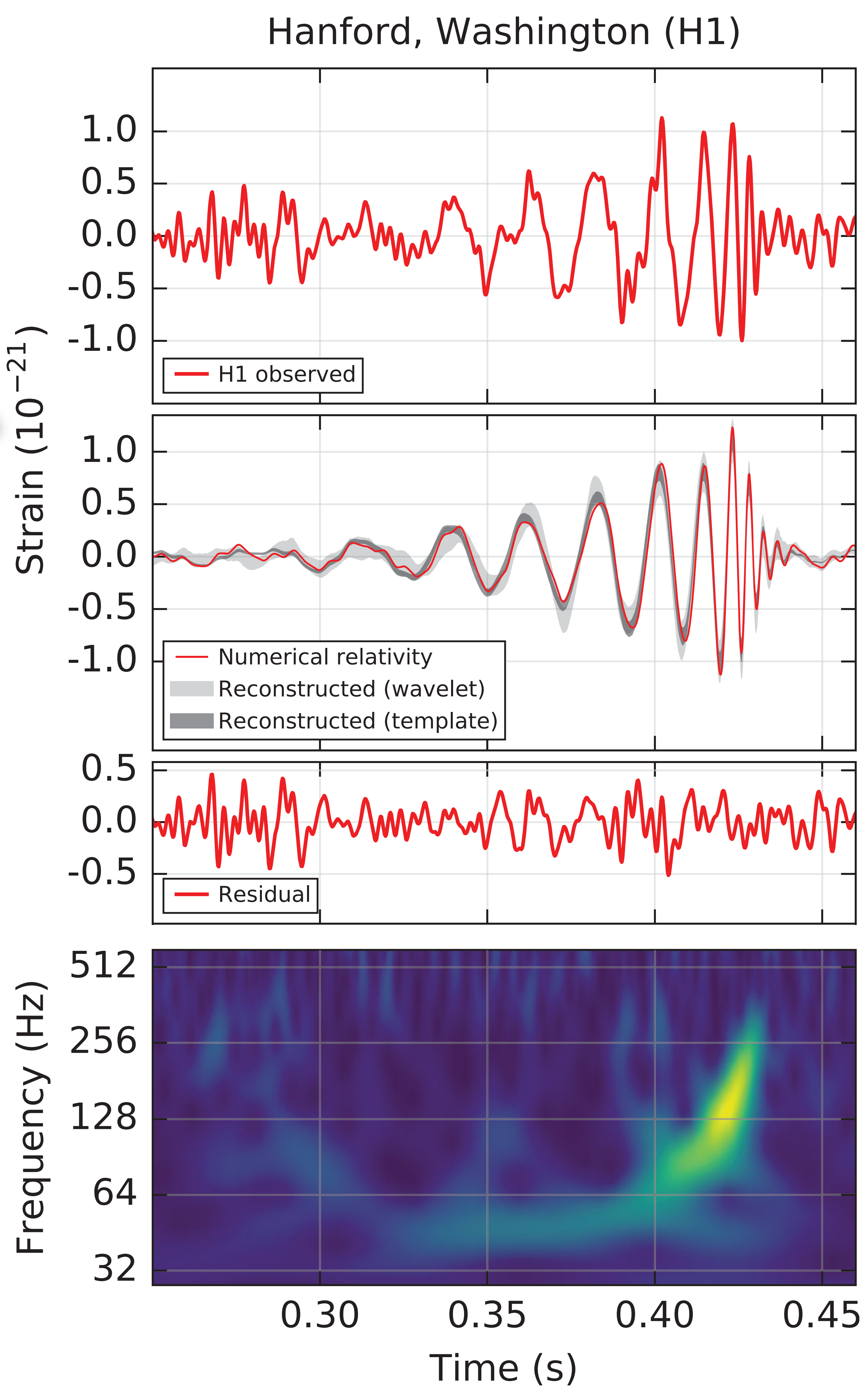
- Observation of gravitational waves by LIGO/Virgo is certainly a breakthrough in fundamental physics
    - First observation on 14th of September 2015:
        - Two ~30 solar mass black holes merging about 1.3 billion light-years from Earth.
    - Frequency range: 10-1000 Hz

See today's talk by Maria Haney

- However, there should/could be many other sources of gravitational waves
    - Primordial black hole merges
    - Boson clouds (BH superradiance)
    - …

See today's talk by Diego Blas

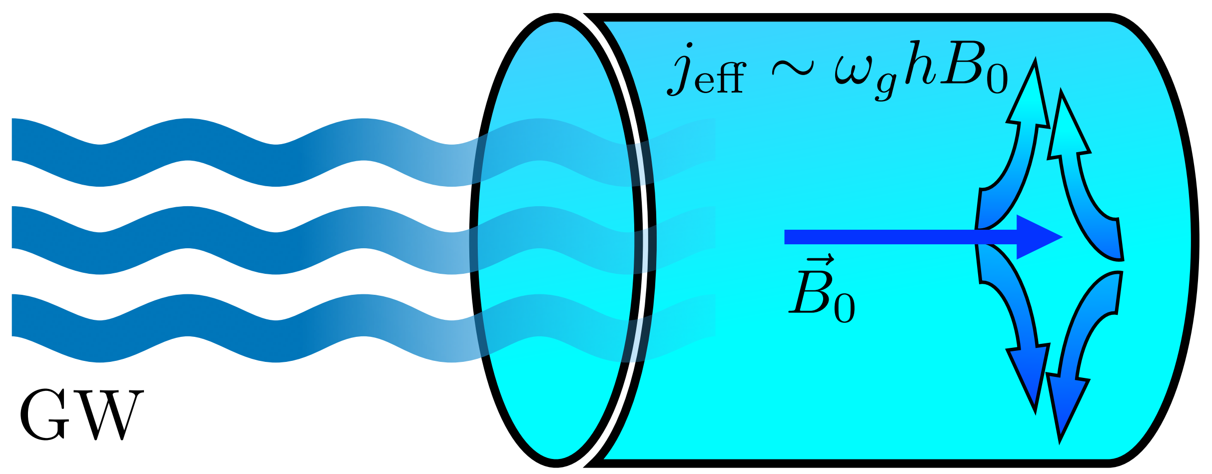
- Those GW would have frequencies in the GHz regime
    - Should search for high frequency GW



[PRL 116, 061102 (2016)]

Conversion of GWs into Photons

[arXiv:2112.11465]



- Two contributing effects
  - Assuming conversion cavity with volume V within const. B-Field

- GW deforms cavity
  - Oscillating change of magnetic flux
  - Excitation of EM field

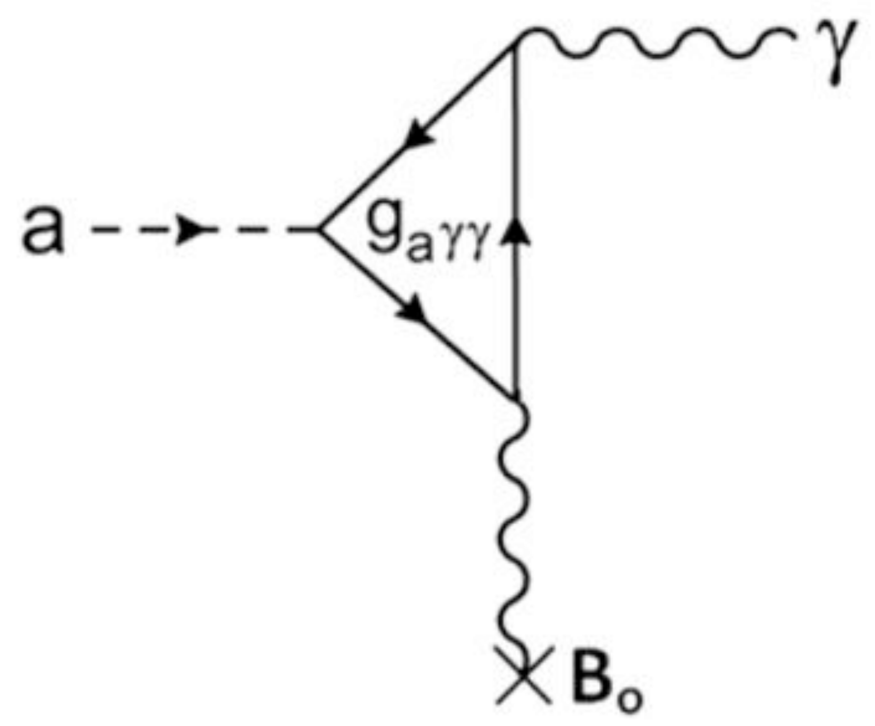
See talk on Monday by Lars Fischer in T01

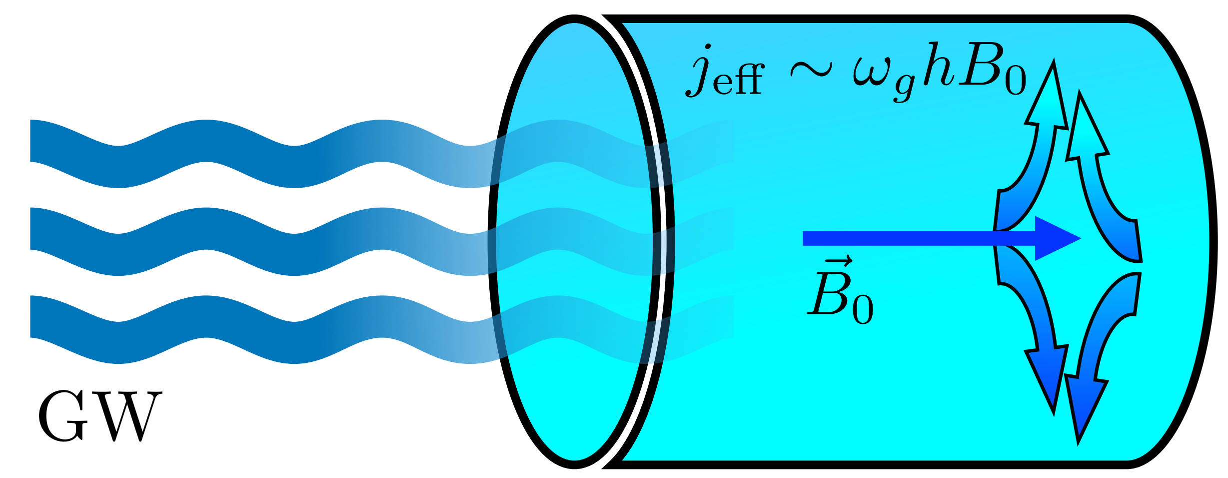
- Direct conversion of gravitons to photons via the inverse Gertsenshtein effect

- Resonant excitation of EM field in Cavity
  - Produced EM power given by:

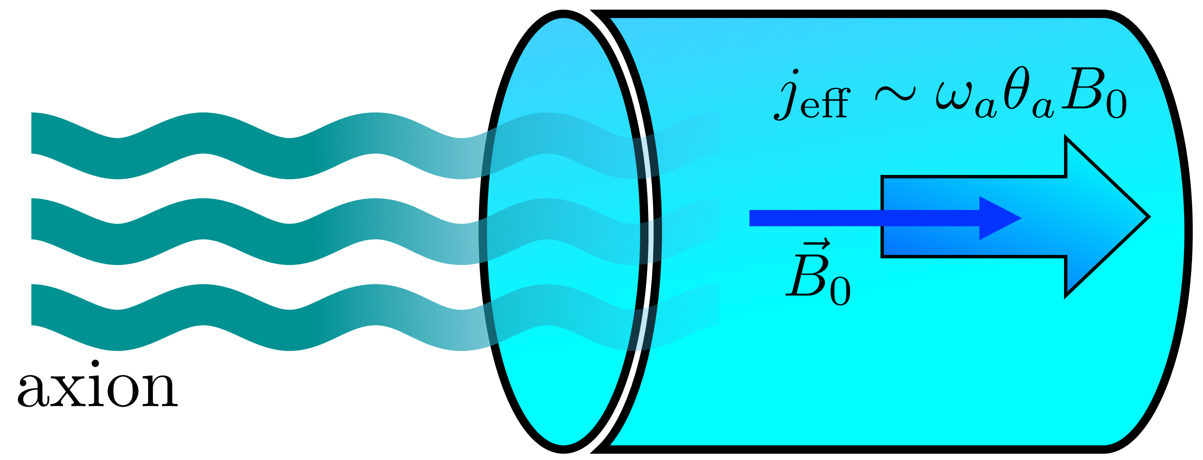
$$P_{sig} = \frac{1}{2} Q \omega_g^3 V^{5/3} (\eta_n h_0 B_0)^2 \frac{1}{\mu_0 c^2}$$

- Axion Haloscopes:
  - RF cavity in magnetic field -> Primakov conversion of axions to photons
  - Resonant excitation of cavity mode
  - **Cavity based haloscopes are sensitive to GWs**





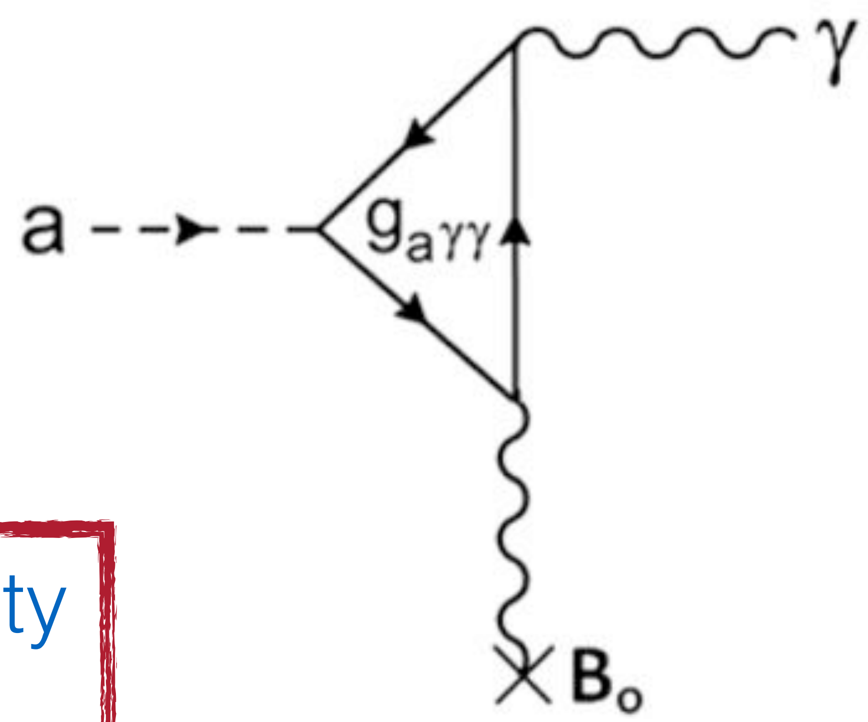
- GW:
  - Typical quadruple structure
  - Preferred mode: TM 020
  - Current direction dependent on GW



- Axions:
  - Preferred mode: TM 010
  - Current dependent on B-field direction
  - Litle overlap with GW mode

[arXiv:2112.11465]

Similarity to Axion Searches



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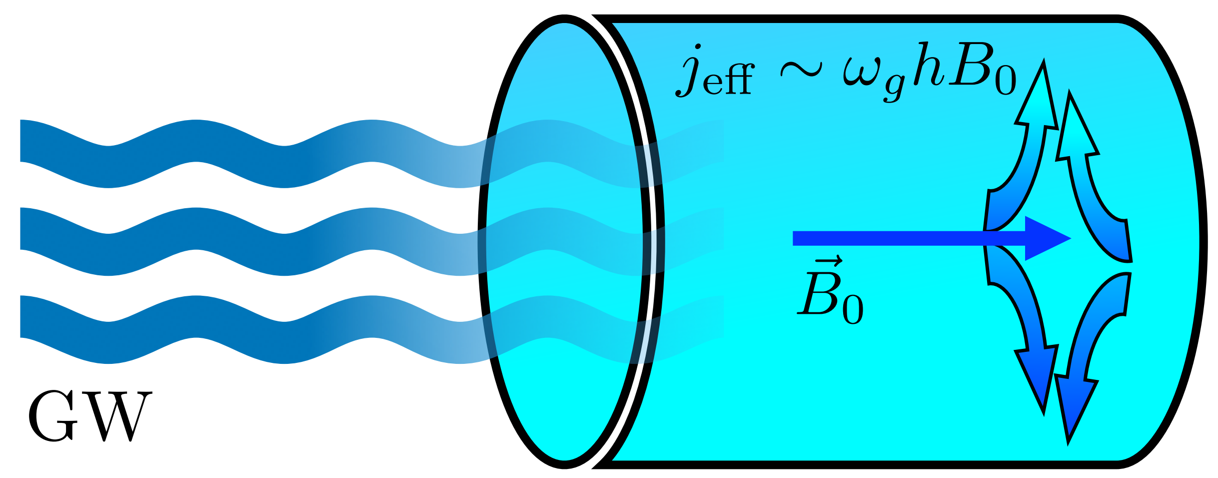
Interesting sensitivity range for PBH



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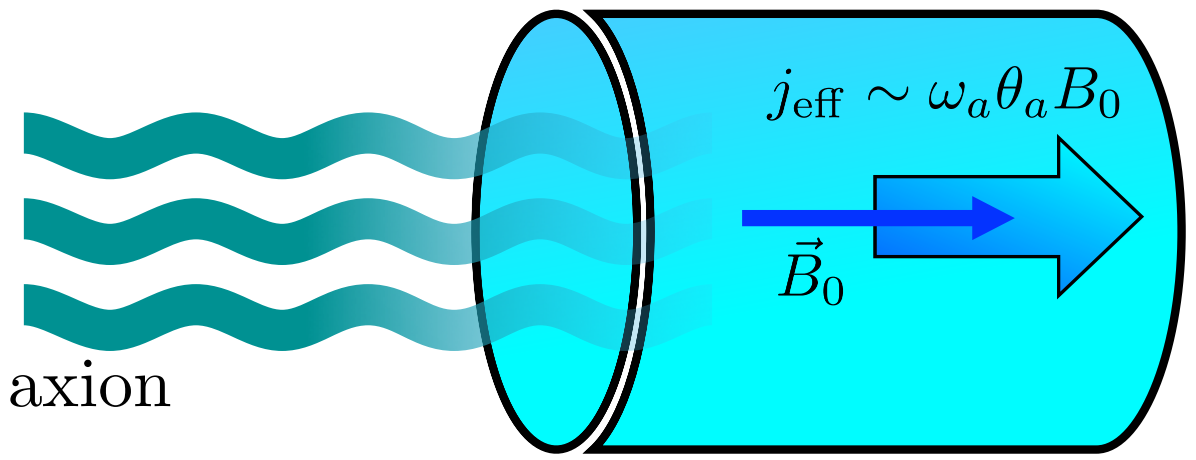

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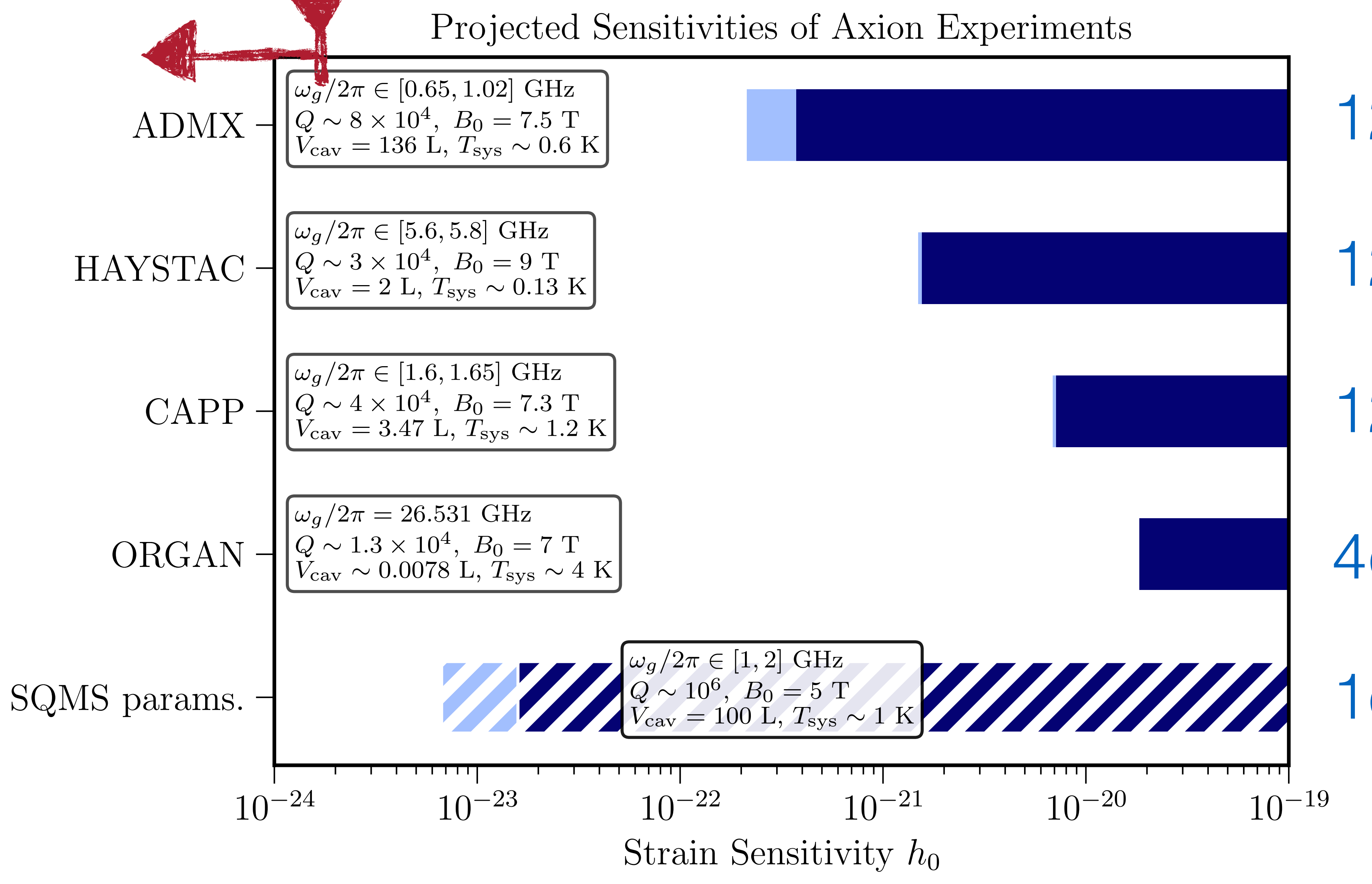
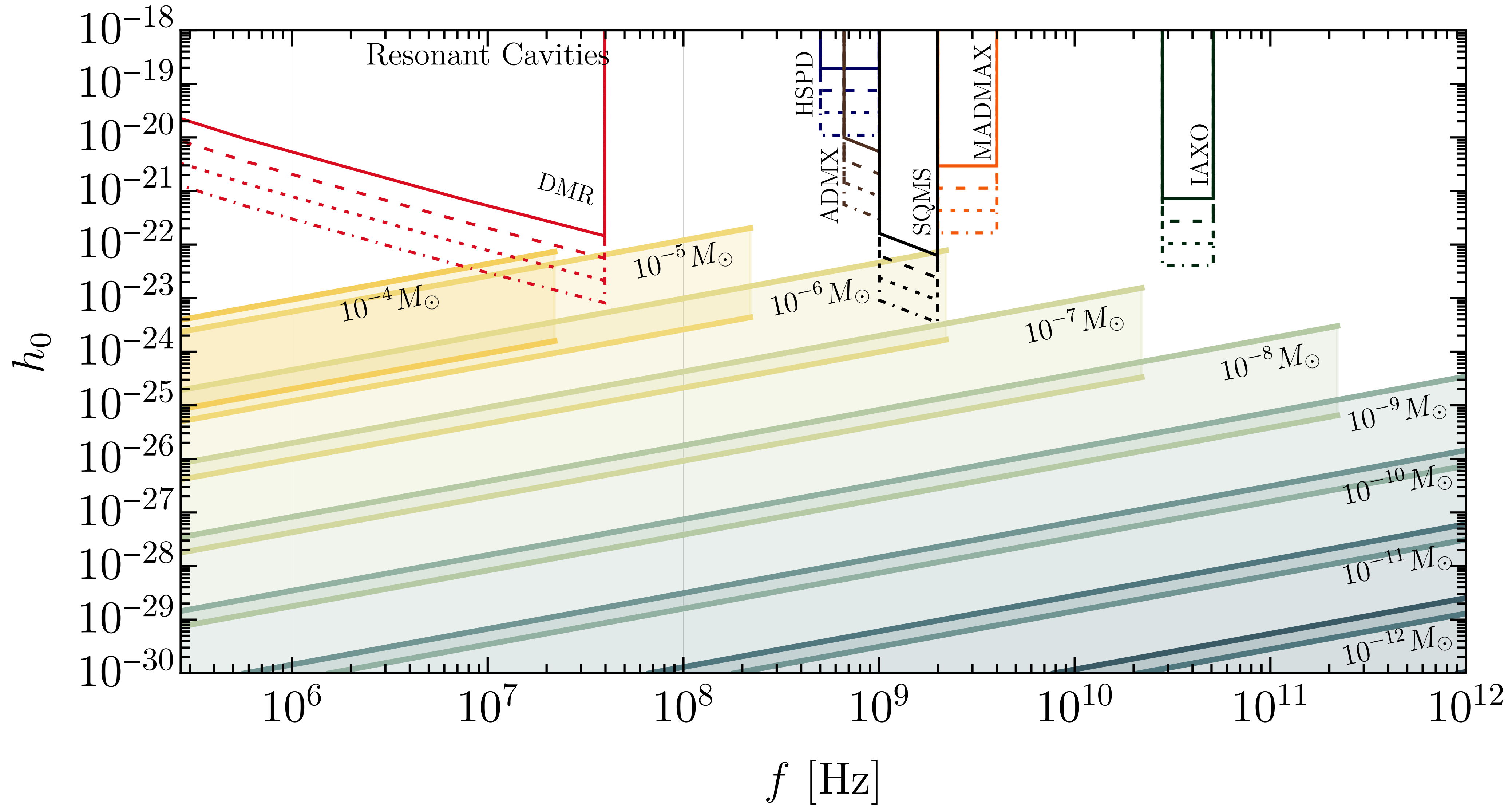


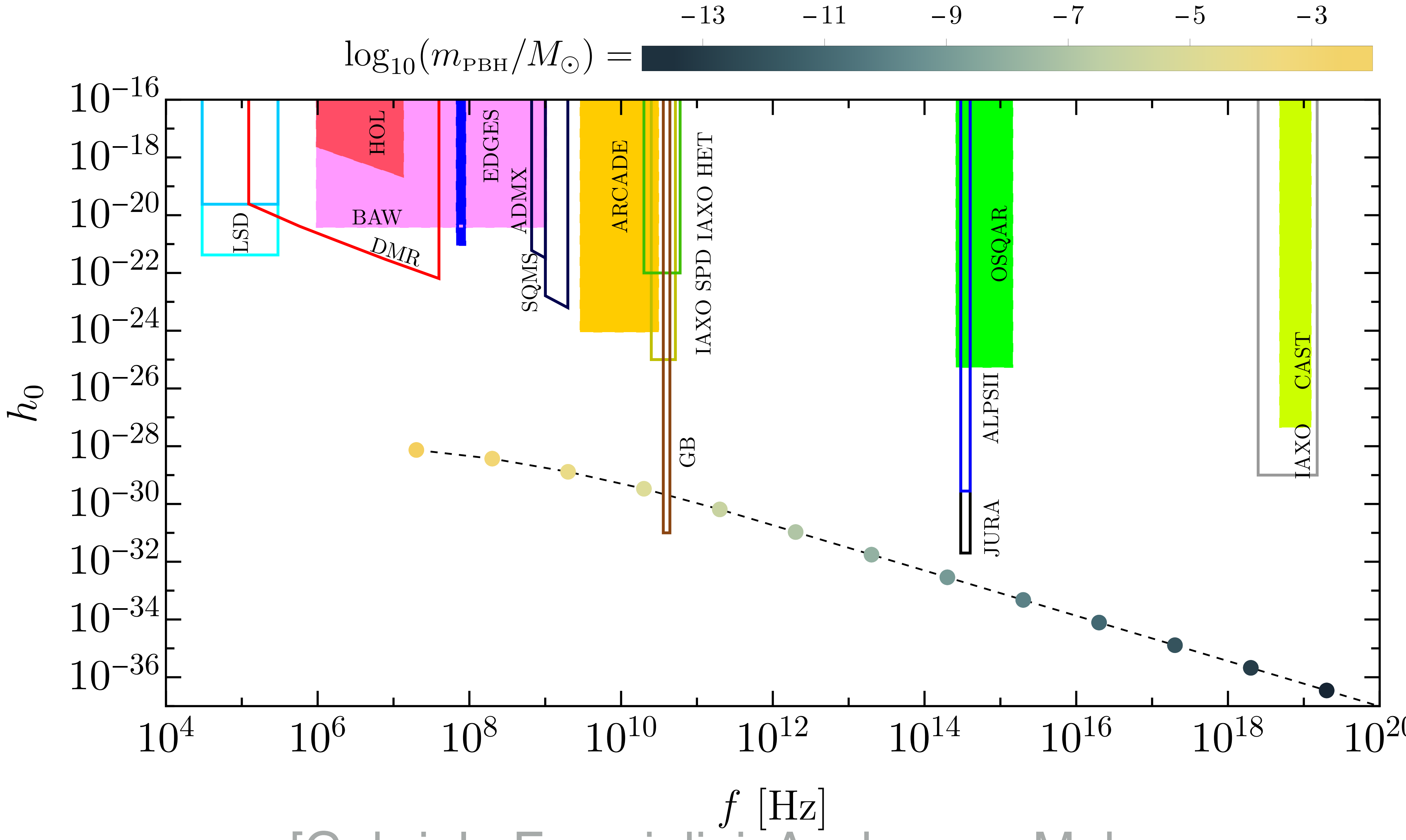
t_int
120s
120s
120s
4d
1d

[arXiv:2112.11465]

[Detecting high-frequency gravitational waves with microwave cavities
Asher Berlin, Diego Blas, Raffaele Tito D'Agnolo , Sebastian A.R. Ellis
arXiv:2112.11465]

- Several well motivated beyond the standard model sources:
  - Primoridal black hole mergers
    - Chirp signals
  - GW from boson clouds around BHs
    - (BH super radiance)
    - Monochromatic over long timescales
  - Stochastic GW background
    - Even lower energy



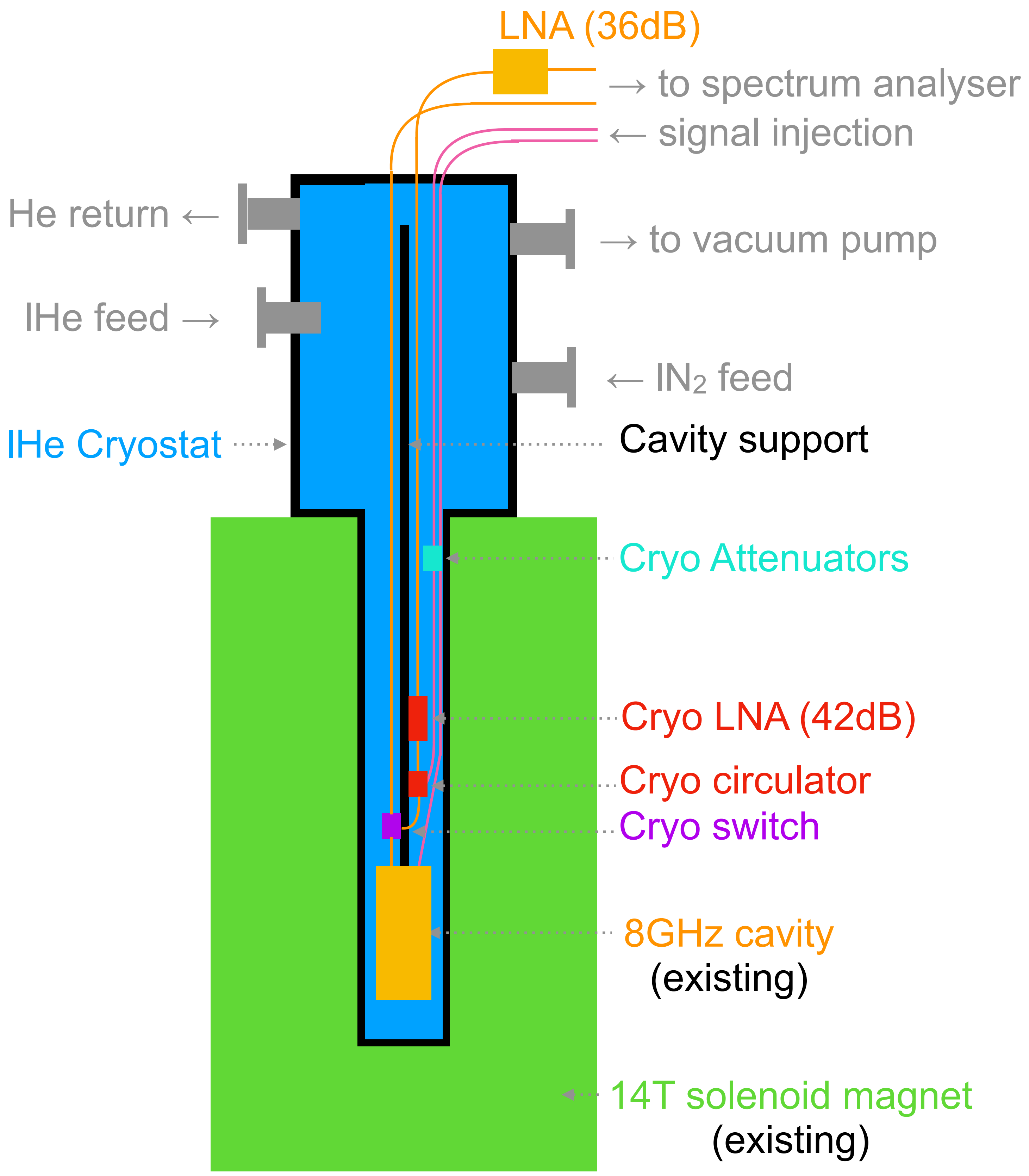


- Displayed expected experimental sensitivities for PBHs:
  - Assuming GW signal long enough to ring up cavity
  - E.g.: given for ADMX,SQMS @ $m_{PBH} \approx 10^{-10} M_\odot$

[Gabriele Franciolini, Anshuman Maharana, Francesco Muia; arXiv:2205.02153v1]

- **Supax:** superconducting axion search @ Mainz
  - First results on dark photons presented in poster by Tim Schneemann
  - Will study spherical cavities
  - Study of new SC materials for resonant cavity experiments

[arXiv:2308.08337]



Metamaterials

$$P_{sig} = \frac{1}{2} Q \omega_g^3 V^{5/3} (\eta_n h_0 B_0)^2 \frac{1}{\mu_0 c^2}$$

- Depends on cavity material:
  - High purity copper: ~$5 \cdot 10^4$
  - Superconducting: difficult in high magnetic field!
    - Target: 10^6
    - Achieved: $3 \cdot 10^5$ (CAPP, non tunable)
      - Materials under study: Nb_3Sn, HTS materials (YBCO)

- Up to 14T magnets in use
  - Up to 20T envisioned
- Larger fields - smaller volume

How to become more sensitive?

- Current efforts focus on **improving single cavity** sensitivity
- But what about **combining various setups**?

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Disclaimer

This is not a fully fledged proposal in all glory detail

Rather intended as basis for discussions

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- Phase aligned combination voltages from of N cavities:

$$V_{comb} = \frac{it\omega}{\sqrt{N}} \sum_i V_i e^{i\phi_i} \propto \sqrt{N} V_0$$

$$V_i = V, \phi_i = \phi$$

- Hence the **signal power scales linearly in N**!
  - Assumed single setup
    - 14T B-field, about 10cm diam., 30cm long
    - 3 spherical cavities @5GHz, SC, high Q
    - 1s integration time

Setup	Supax	GravNet
Shape	cyl.	spher.
f_0 [GHz]	8.3	5.0
Volume [l]	0.128	0.21
Q_0	39600	10^6
η	0.08	0.6
T_{sys} [K]	5	0.1
B [T]	14	
int. time	1 s	
n cavities	1	3
noise power [W]	$1.5 \cdot 10^{-21}\,W$	$6.2 \cdot 10^{-23}\,W$
$h_0(P_{sig} = P_{noise})$	$7.1 \cdot 10^{-21}$	$5.2 \cdot 10^{-23}$

GravNet - a global network for HFWG detection

- How sensitive can we get with 10 setups, scattered around the globe

- Assumptions:
  - Sampling of Waveform -> offline combination of phase aligned IQ data
  - 10 setups as shown before
    - Effective signal power increased by factor 10
    - Strain sensitivity increased by factor $\sqrt{10} \approx 3$

$h_0 < 10^{-23}$, 1 second integration time

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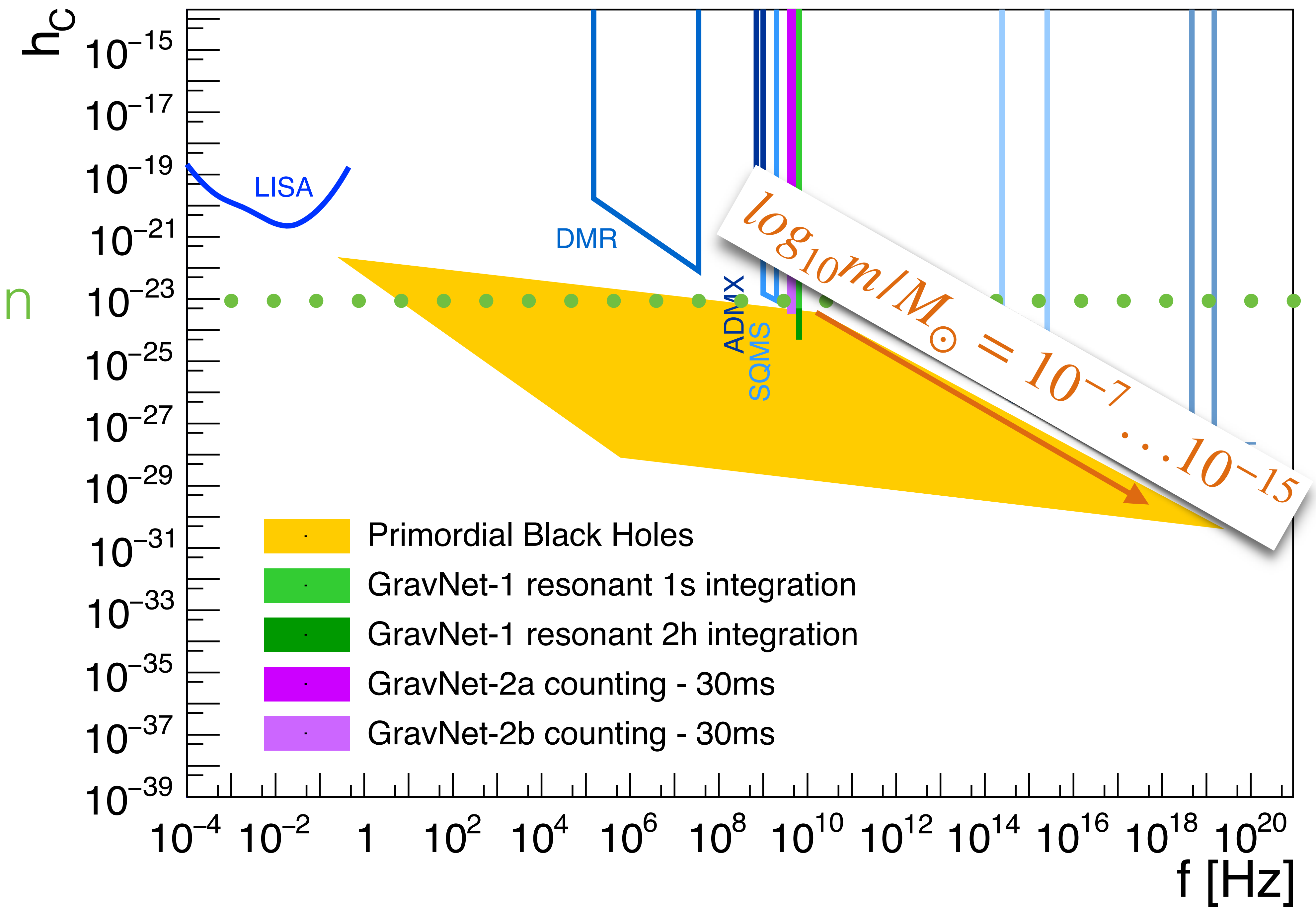
1s integration



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  - Otherwise: scan through all possible directions and repeat combinations

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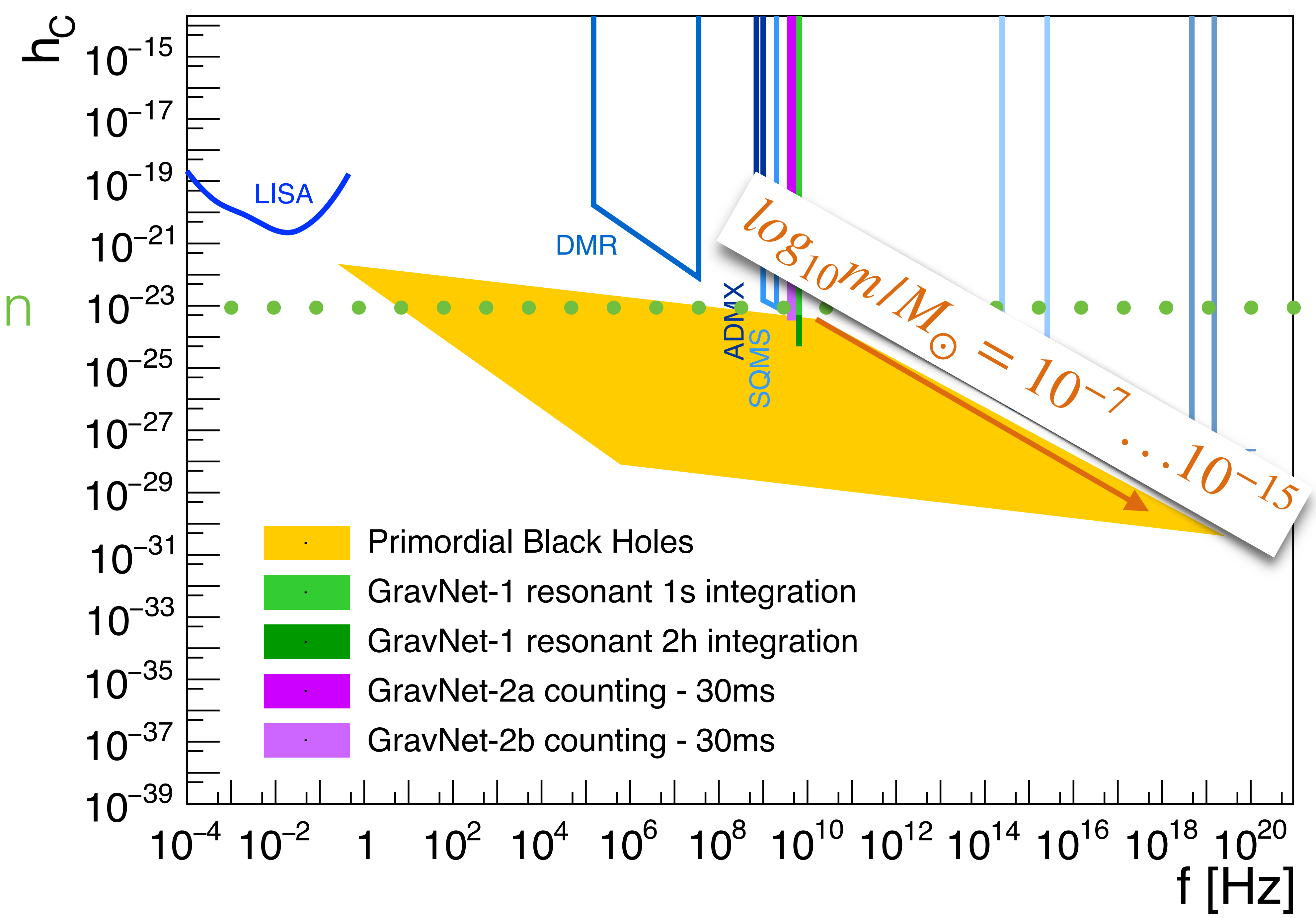
- No frequency tuning needed:
  - PBH signals are fast transients
    - Single frequency sufficiency

1s integration



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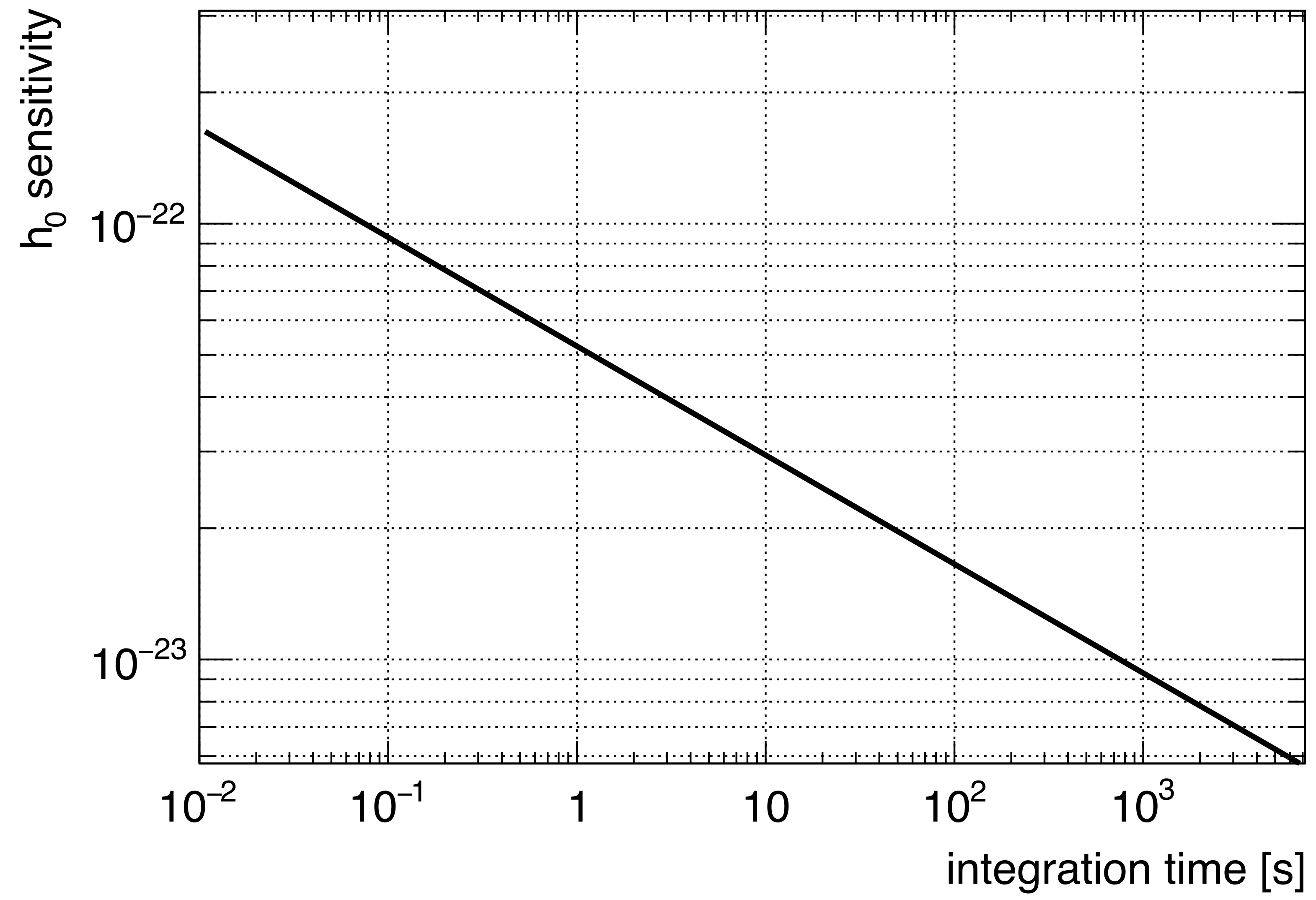
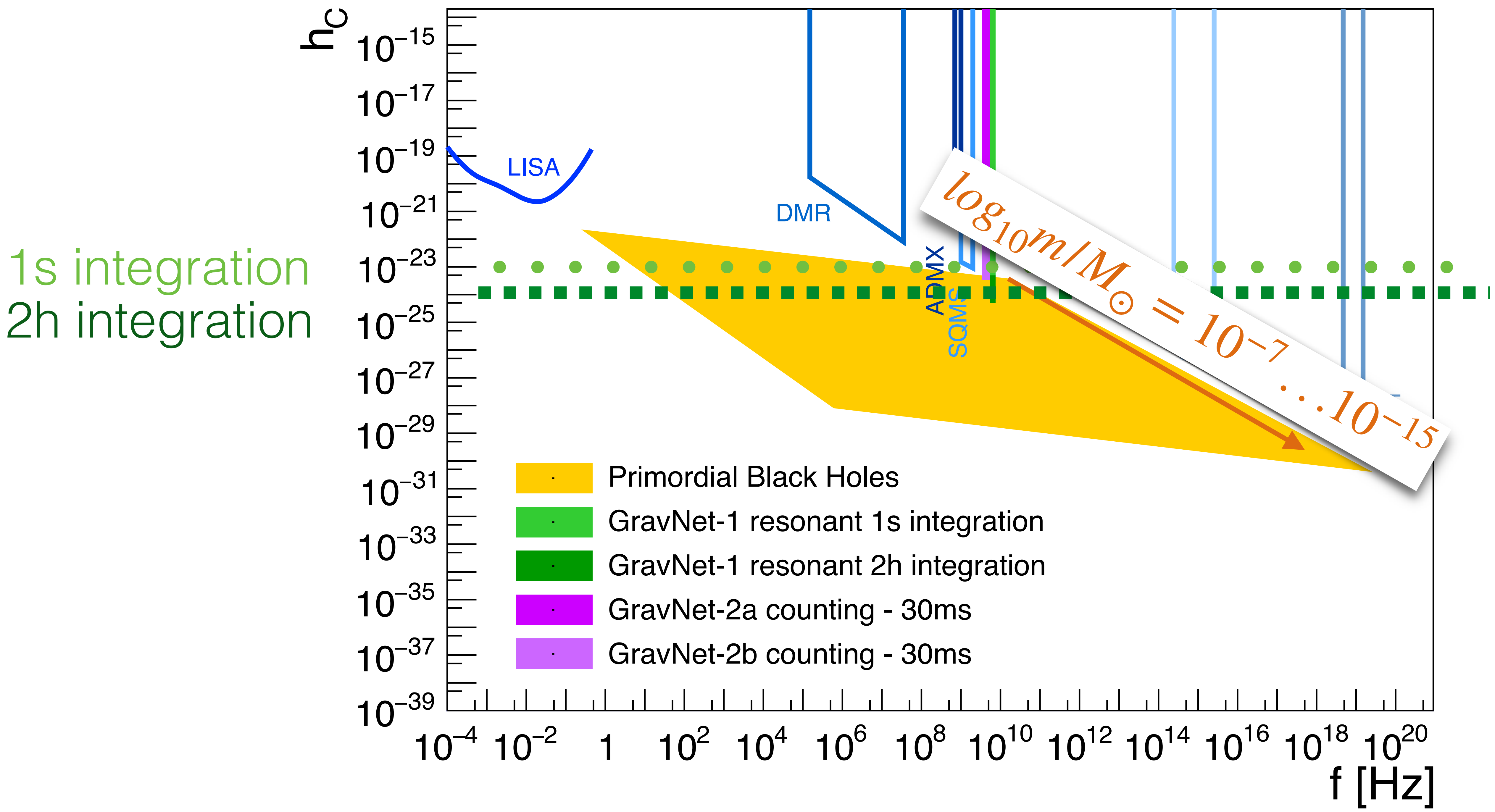
$$h_0 < 10^{-23}\text{, 1 second integration time}$$

- Longer integration times
  - Sensitivity gain with integration time $t^{1/4}$

$$h_0 < 10^{-24}\text{, 2h integration time}$$

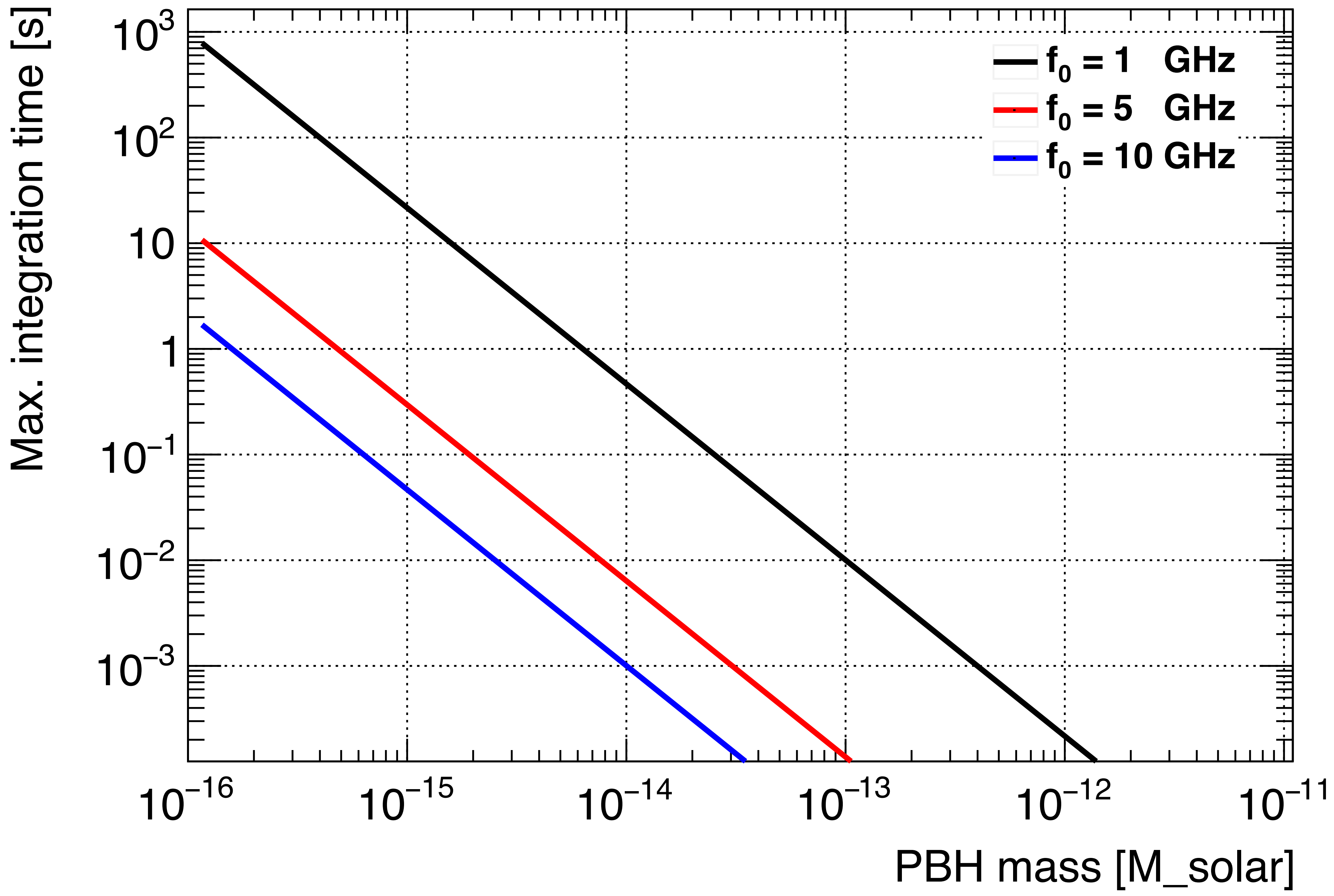




PBH merger & high Q cavities

- GW strain: largest if merging is imminent (closest to innermost stable circular orbit)
  - Frequency drift large

$$\dot{f} = \frac{96}{5}\pi^{8/3} m_c^{5/3} f^{11/3} \simeq 4.62 \times 10^{11}\,\mathrm{Hz}^2 \left(\frac{m_{\mathrm{PBH}}}{10^{-9} M_\odot}\right)^{5/3} \left(\frac{f}{\mathrm{GHz}}\right)^{11/3}$$

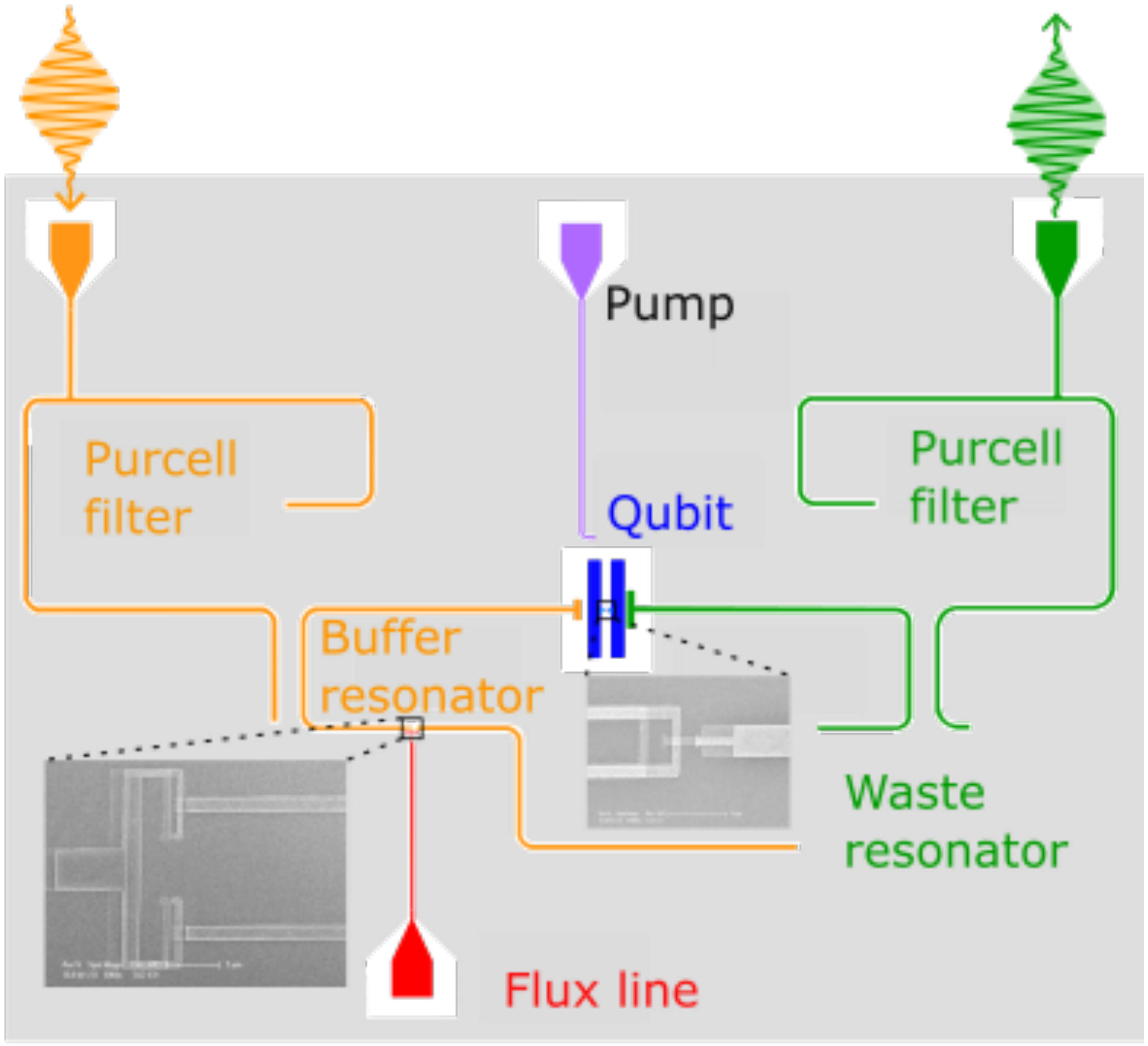


- To resonantly excite a cavity:
  - GW frequency must stay within resonator bandwidth
    - $\omega/Q \approx 10^{10} Hz/10^6 = 10 kHz$
  - Very short integration times O(ms) or below for larger BH masses
- Resonant detection difficult for HFGWs from PBH mergers!
  - Alternative?

Alternative for fast transient signals - Photon Counting

- Recent progress in R&D for single RF photon counters

- Several technologies under study
  - Current Biased Josephson Junctions [arXiv:2302.07556]
  - Kerr Josephon Parametric amplifiers [arXiv:2308.07084]
  - Transmon Q-Bit readout [arXiv:2307.03614]



[arXiv:2307.03614]

- Shown single photon efficiency: 43% @ 90 Hz dark count rate

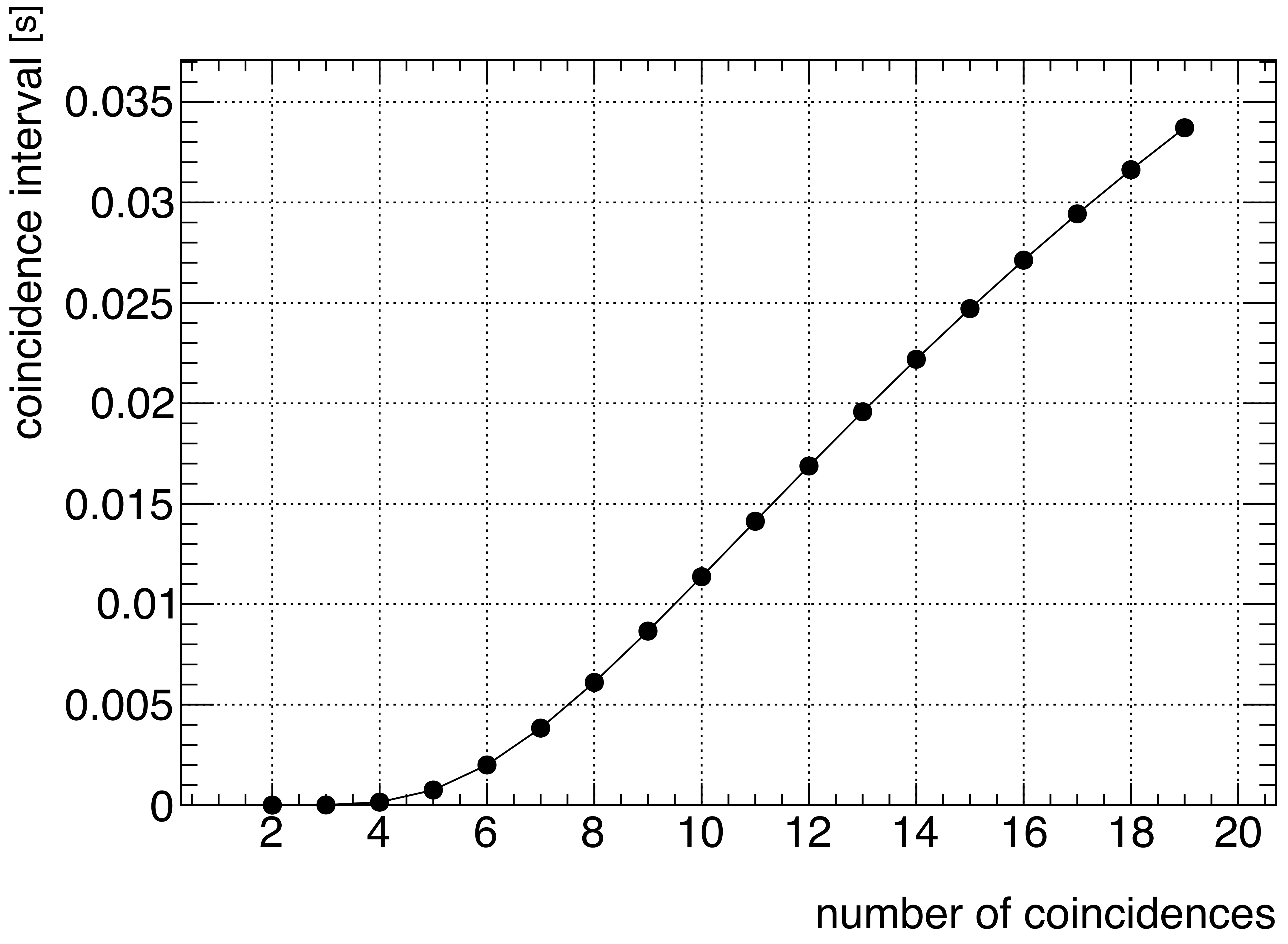
- Measurement boils down to a **coincidence measurement** !

- Coincidence window and needed number of coincident detectors optimised depending on
  - Background rate (thermal, detector noise)
  - Signal Rate

Photon Counting - Coincidence Interval

- Background rate:
  - Average thermal power in cavity @ 0.1K ~ 4×10^{-23} W, corresponding to 10 photons / s @ 5 GHz
  - Could be lowered going to lower temperatures (challenging)
  - Assuming advances in the near future on the single photon sensors:
    - Detector dark count rate will drop significantly -> negligible

- Parameter used for Calculation:
  - Allowed accidental coincidence rate: <= 1/year
  - Background rate: 10 Hz
  - N detectors: 20



Photon Counting - Signal efficiency

- Overall signal efficiency dependent on detector efficiency, coincidence window and signal photon flux:
  - $\epsilon_{single} = \epsilon_{det} \Delta t_{coincidence} \Phi_{sig}$ Φ_{sig} = signal photon flux
  - $\epsilon_{tot} = \sum_{i>k} \binom{N}{k}, p = \epsilon_{single}$, k = number of required coincidences, N = number of detectors

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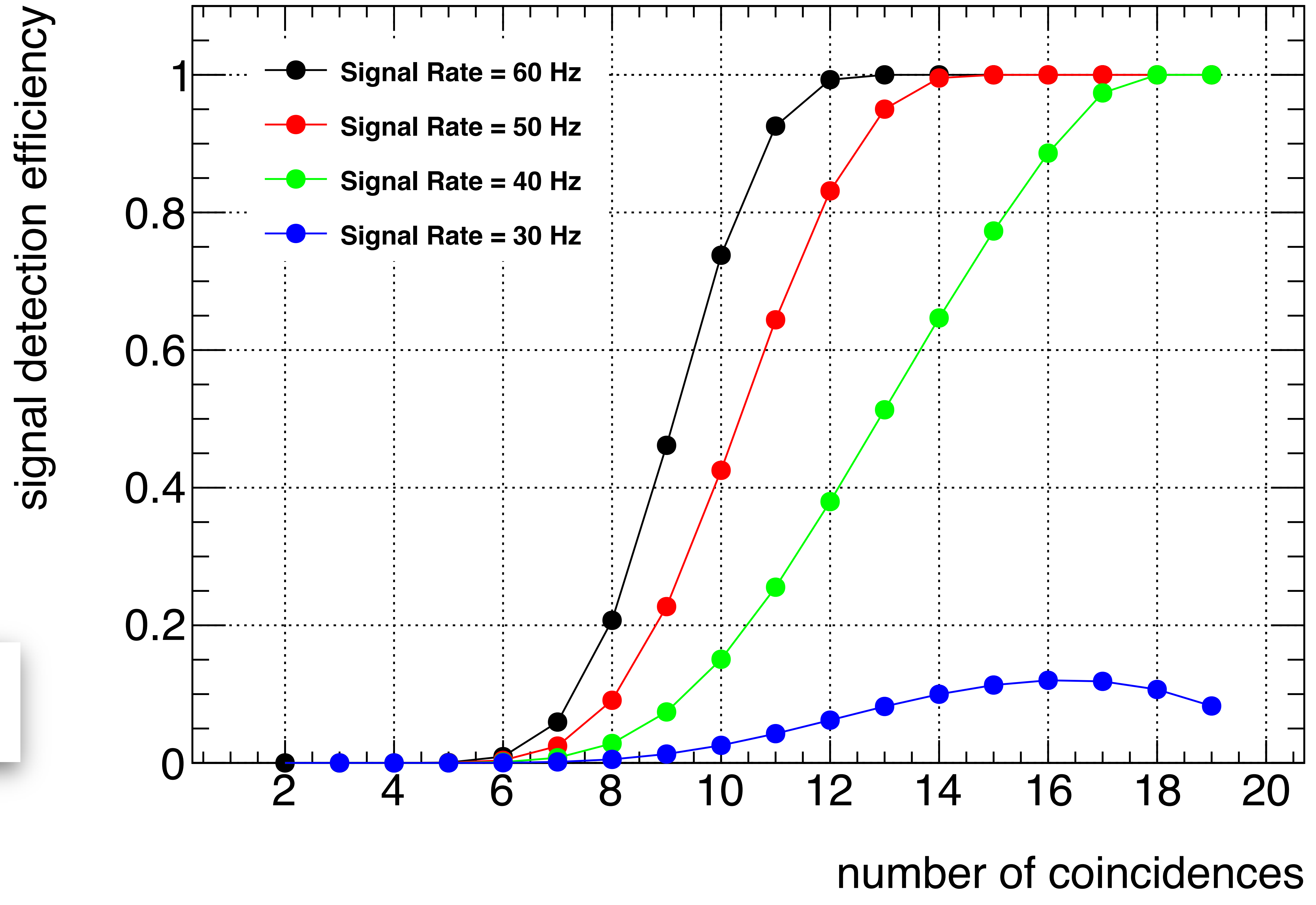

- With **20 detectors** a photon flux of **40 Hz** can be detected with an efficiency of 1 within a coincidence interval of **32ms**

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- Signal photon flux depends on conversion region:
  - a) Magnet dimensions as before (9cm diameter), B = 14T
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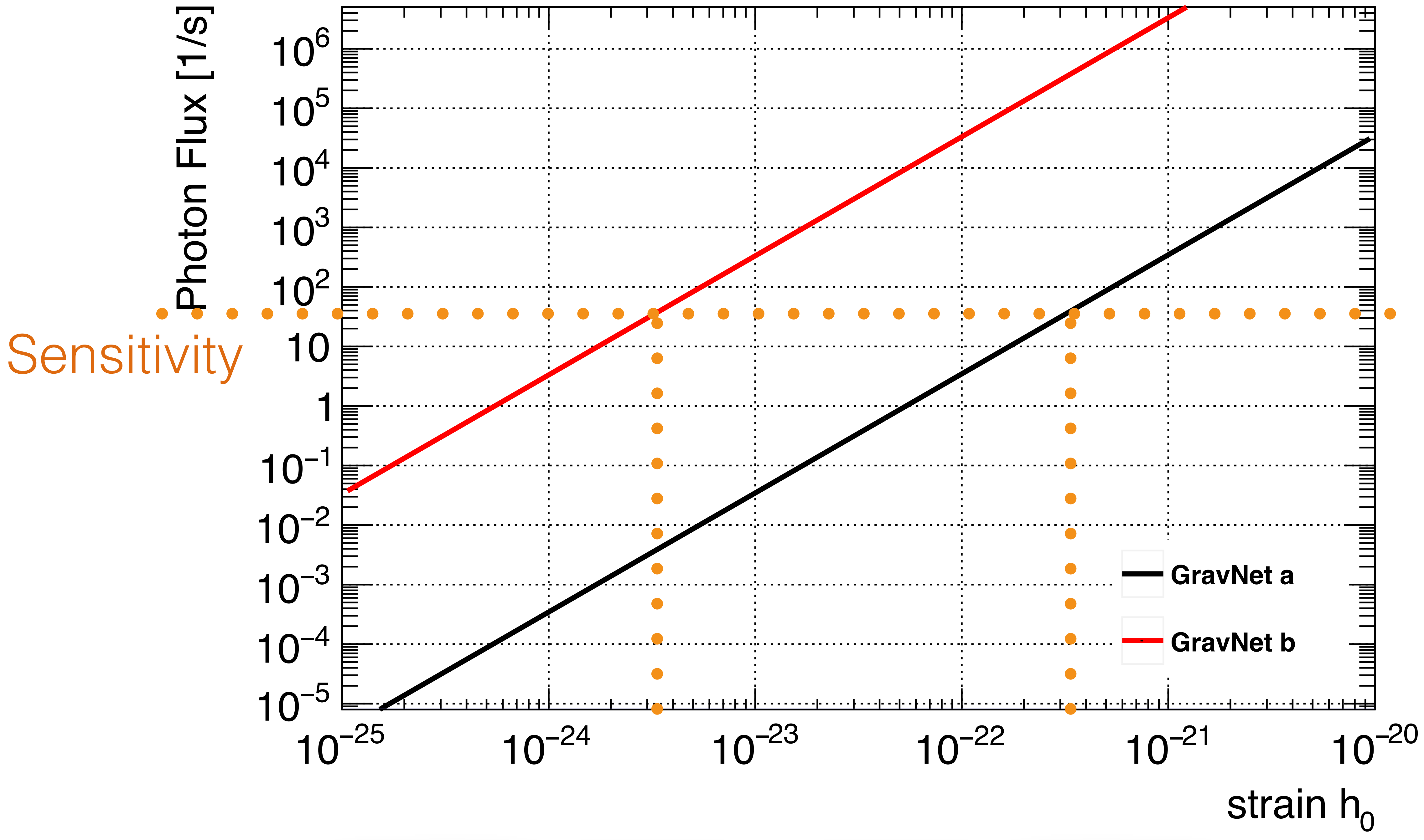
Setup	GravNet-a	GravNet-b
radius	40 mm	40 cm
length	12cm	50 cm
Volume $[m^3]$	6×10^{-4}	0.25
Q_0	10^6	10^5
$T_{\rm sys}$ [K]	0.1	0.1
B [T]	14	9



- Achievable sensitivity:
  - **$h_0 < 3\text{x}10^{-22}$ .... $3\text{x}10^{-24}$**
  - With coincidence time of 32ms!

Photon Counting - Sensitivity on GW

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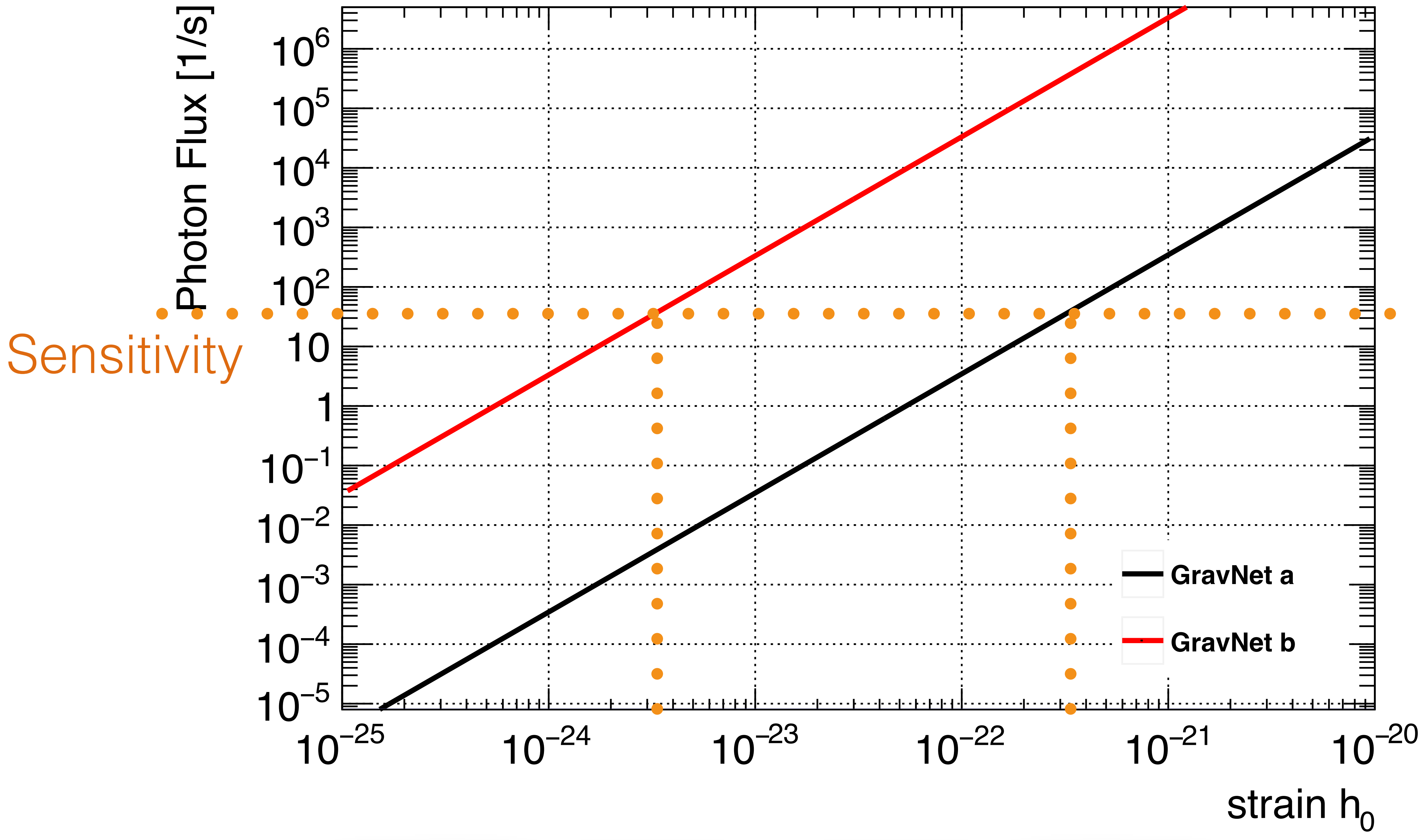
Global network of HFGW detectors will be able to reach into the interesting region for PBH with existing technologies!



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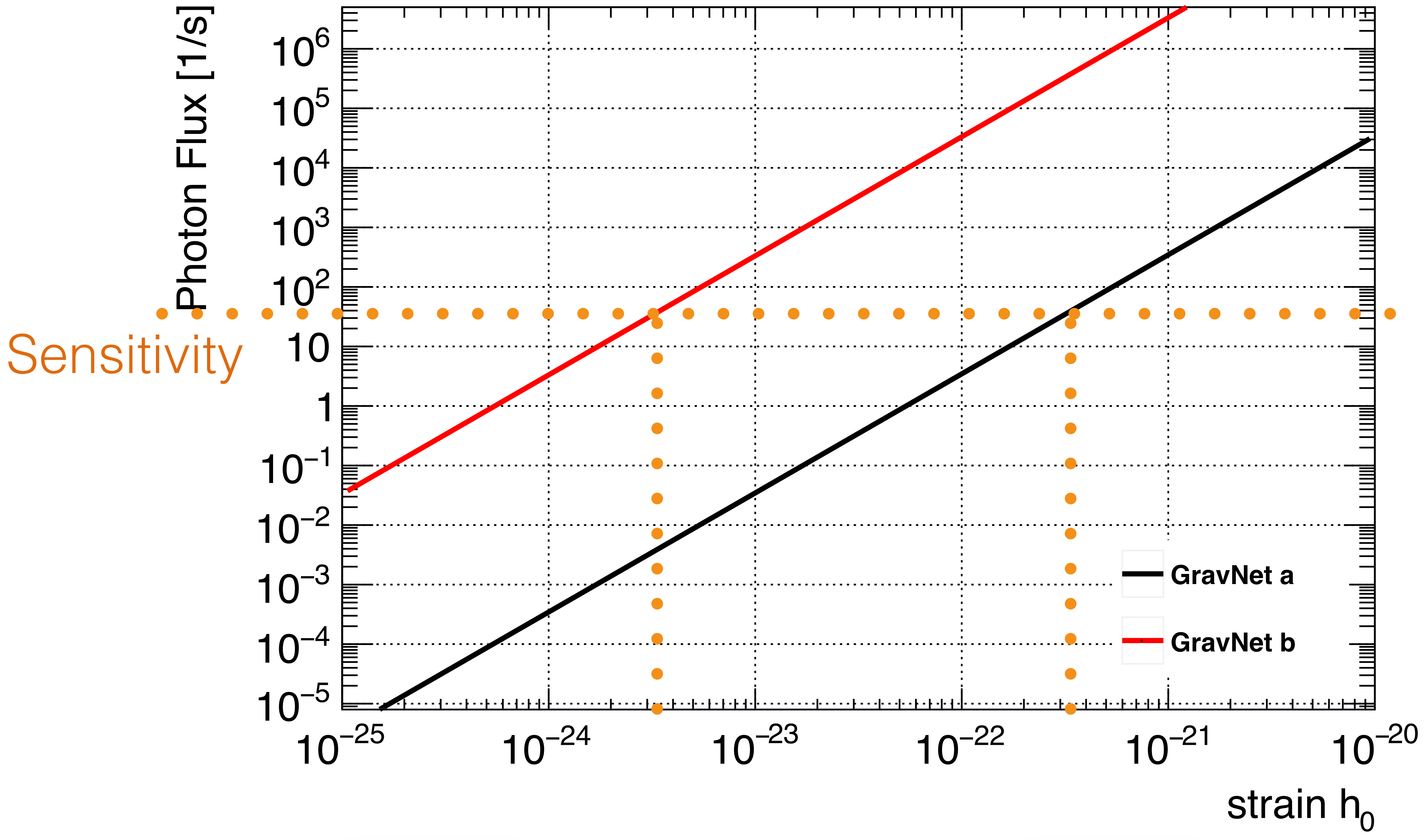

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  - **$h_0 < 3x10^{-22}$ .... $3x10^{-24}$**
  - With coincidence time of 32ms!

Significant room for improvements:
more detectors, larger volumes, higher detector efficiency

Summary

- To increase the sensitivity of halo scope style experiments we suggest to build a **global network of detectors**
  - Remember: SNR scales linear with number of detectors!

- Integrating measurement:
  - Sample RF data and combine phase aligned

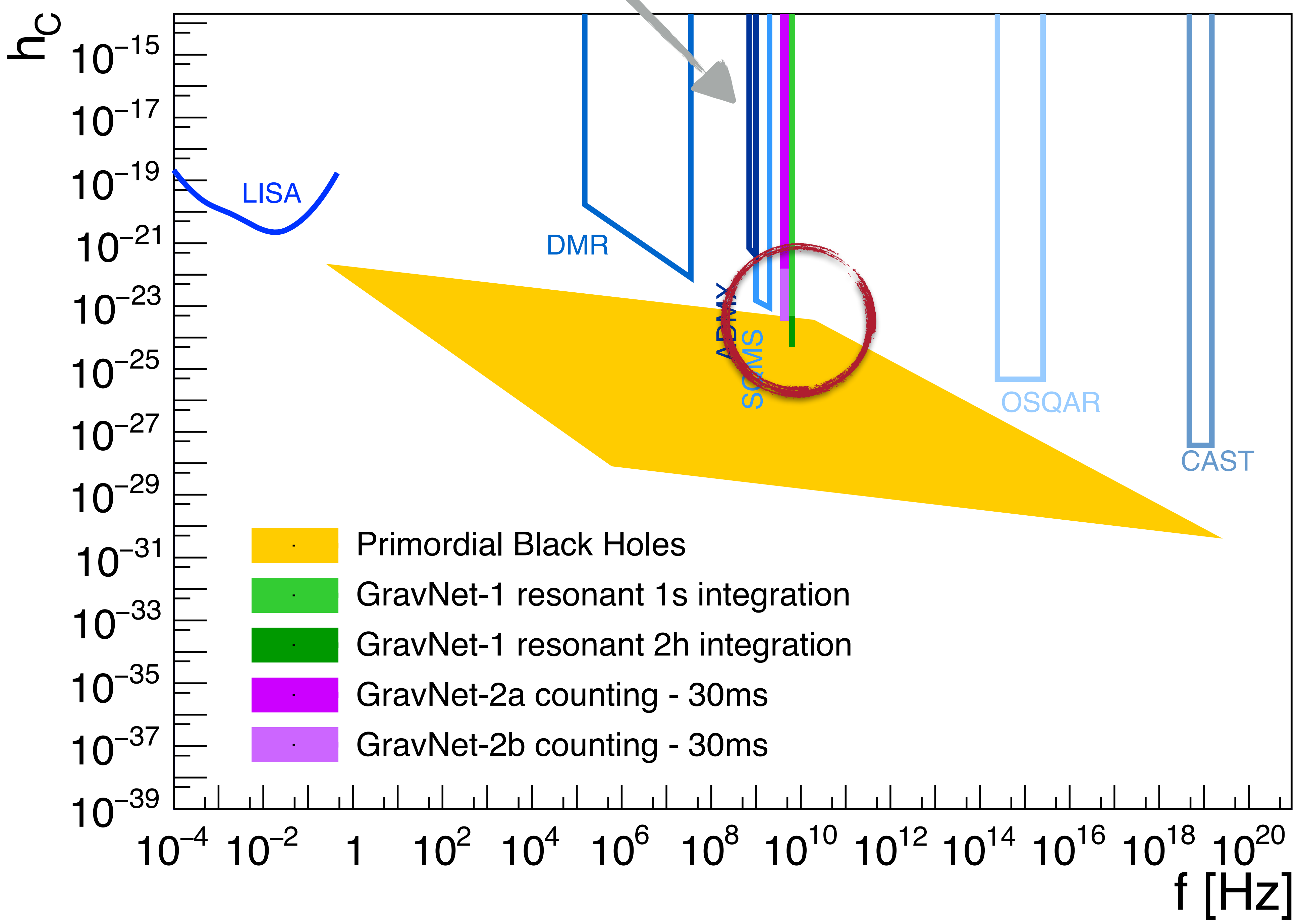
- Typical integration times too long to be sensitive to BH merges!

- Photon counting style experiments:
  - Recent advancements in single RF photon detection allows to use coincidences of several detectors
  - Using 20 independent detectors:
    - Sensitivity: $h_0 < 3x10^{-22}$ …. $3x10^{-24}$

Requires large meta material cavities
(high frequency @ large volume)

- Single frequency sufficient to hunt for PBH mergers!
- Could even combine measurements at different frequencies

t_{int} > O(100s) [arXiv:2112.11465]



[arXiv:2308.11497]

Conclusion

GravNet is an idea up for discussion

- Many advantages in combining efforts to hunt for UHFGWs in coordinated way
- GravNet would significantly improve the sensitivity on high frequency gravitational waves
  - Based on commercial magnet systems, which is comparatively cheap
  - Worldwide collaboration would share costs automatically with local lab-based experiments
  - Easy exchange of R&D results and integration at all locations
- Sensitivity to the PBH parameter space with existing technologies!

Interested? Let's have a coffee!