Andreas Ringwald Fest

Z. Fodor Penn State, Univ. Wuppertal, FZ Juelich, Univ. Budapest, UCSD

Phys. Rev. Lett. 88 (2002) 171101 & Nature 539 (2016) 7627

September 27 2024, DESY, Hamburg, Germany



• 12 publications with 7 original papers

- 12 publications with 7 original papers
- from 2001 to 2004 ultra high energy cosmic rays (UHECR)

- 12 publications with 7 original papers
- from 2001 to 2004 ultra high energy cosmic rays (UHECR)
- from 2015 to 2016 axion physics

- 12 publications with 7 original papers
- from 2001 to 2004 ultra high energy cosmic rays (UHECR)
- from 2015 to 2016 axion physics

It was really a great pleasure to work with Andreas! Thank you!

- 12 publications with 7 original papers
- from 2001 to 2004 ultra high energy cosmic rays (UHECR)
- from 2015 to 2016 axion physics

It was really a great pleasure to work with Andreas! Thank you!

We wanted to intensify/institutionalize \Rightarrow excellence initiative

- 12 publications with 7 original papers
- from 2001 to 2004 ultra high energy cosmic rays (UHECR)
- from 2015 to 2016 axion physics

It was really a great pleasure to work with Andreas! Thank you!

We wanted to intensify/institutionalize \Rightarrow excellence initiative

I will choose one work from each time period

- Z-burst explanation of UHECRs and neutrino masses
- topological susceptibility at large temperatures

```
ultra-high-energy cosmic ray (UHECR): energy greater than 1 EeV (10^{18} \text{ eV})
```

```
ultra-high-energy cosmic ray (UHECR): energy greater than 1 EeV (10^{18} eV) University of Utah's Fly's Eye 1991: 3.2 \cdot 10^{21}eV nucleus as a 100 km/h baseball
```

```
ultra-high-energy cosmic ray (UHECR): energy greater than 1 EeV (10<sup>18</sup> eV)

University of Utah's Fly's Eye 1991:
3.2·10<sup>21</sup>eV nucleus as a 100 km/h baseball

GZK cutoff-paradox: why do we observe them? within 50 Mpc they should interact with CMB
```

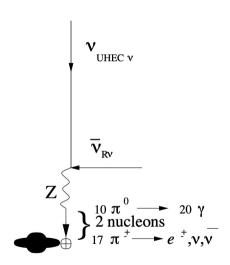
One should see a cutoff somewhere below 10²⁰ eV

ultra-high-energy cosmic ray (UHECR): energy greater than 1 EeV (10^{18} eV)

University of Utah's Fly's Eye 1991: 3.2·10²¹eV nucleus as a 100 km/h baseball

GZK cutoff-paradox: why do we observe them? within 50 Mpc they should interact with CMB One should see a cutoff somewhere below 10^{20} eV

UHECR neutrinos hitting relic neutrinos Z-burst ⇒ UHECR hadrons scenario depends on the neutrino mass



Relic neutrinos: decoupled at T=1 MeV and t=1 second

Relic neutrinos: decoupled at T=1 MeV and t=1 second

Their predicted number density: 56/cm³ (comparable to CMB)

Relic neutrinos: decoupled at T=1 MeV and t=1 second

Their predicted number density: 56/cm³ (comparable to CMB)

Z resonance production: $\nu \bar{\nu} \longrightarrow Z \longrightarrow \text{hadrons}$

Relic neutrinos: decoupled at T=1 MeV and t=1 second

Their predicted number density: 56/cm³ (comparable to CMB)

Z resonance production: $\nu\bar{\nu}\longrightarrow Z\longrightarrow$ hadrons

$$E_{
u}^{
m res} = M_Z^2/2m_{
u} = 4.2 \cdot 10^{21} \; {
m eV} \cdot \! \left(1 \; {
m eV}/m_{
u}
ight)$$

Relic neutrinos: decoupled at T=1 MeV and t=1 second

Their predicted number density: 56/cm³ (comparable to CMB)

Z resonance production: $\nu\bar{\nu} \longrightarrow Z \longrightarrow \text{hadrons}$

$$E_{
u}^{res} = M_Z^2/2m_{
u} = 4.2 \cdot 10^{21} \; \mathrm{eV} \cdot \! \left(1 \; \mathrm{eV}/m_{
u}
ight)$$

decreasing u flux \Rightarrow largest $m_{
u}$ needs the smallest E: dominant

Relic neutrinos: decoupled at T=1 MeV and t=1 second

Their predicted number density: 56/cm³ (comparable to CMB)

Z resonance production: $\nu \bar{\nu} \longrightarrow Z \longrightarrow$ hadrons

$$E_{
u}^{res} = M_Z^2/2m_{
u} = 4.2 \cdot 10^{21} \; \mathrm{eV} \cdot (1 \; \mathrm{eV}/m_{
u})$$

decreasing u flux \Rightarrow largest $m_{
u}$ needs the smallest E: dominant

We can fit m_{ν} , absolute scale, to the observed UHECR spectrum Result: 0.08 eV $< m_{\nu} <$ 0.40 eV, similar scale as the mass differences

Relic neutrinos: decoupled at T=1 MeV and t=1 second

Their predicted number density: 56/cm³ (comparable to CMB)

Z resonance production: $\nu \bar{\nu} \longrightarrow Z \longrightarrow$ hadrons

$$E_{
u}^{
m res} = M_Z^2/2m_{
u} = 4.2 \cdot 10^{21} \; {
m eV} \cdot (1 \; {
m eV}/m_{
u})$$

decreasing u flux \Rightarrow largest $m_{
u}$ needs the smallest E: dominant

We can fit m_{ν} , absolute scale, to the observed UHECR spectrum Result: 0.08 eV $< m_{\nu} <$ 0.40 eV, similar scale as the mass differences

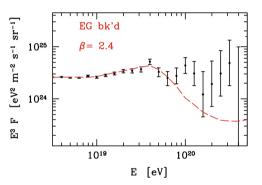
After over two decades, we still don't know the absolute scale

At least three great successes

- (a) explain UHECR & the GZK cutoff-paradox
- (b) give an absolute value for $m_{
 u}$
- (c) prove the existence of relic neutrinos

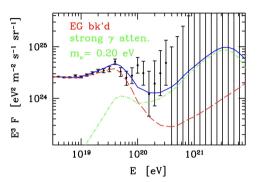
At least three great successes

- (a) explain UHECR & the GZK cutoff-paradox
- (b) give an absolute value for $m_{
 u}$
- (c) prove the existence of relic neutrinos



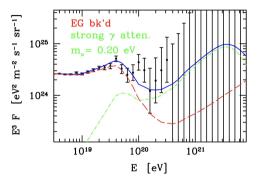
At least three great successes

- (a) explain UHECR & the GZK cutoff-paradox
- (b) give an absolute value for $m_{
 u}$
- (c) prove the existence of relic neutrinos



At least three great successes

- (a) explain UHECR & the GZK cutoff-paradox
- (b) give an absolute value for m_{ν}
- (c) prove the existence of relic neutrinos



2017: instead of the 30 UHECR events, what could have confirmed AGASA Auger saw only two events \Rightarrow GZK cutoff is there (Earlier large E/flux results: probably incorrect energy calibration issue)



differential proton flux is proportional to the density of relic neutrinos

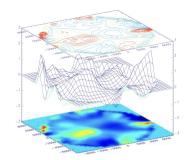
differential proton flux is proportional to the density of relic neutrinos it follows the total mass distribution however, with less clustering

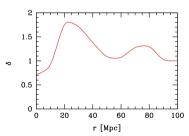
differential proton flux is proportional to the density of relic neutrinos it follows the total mass distribution however, with less clustering vary the density between these two cases

differential proton flux is proportional to the density of relic neutrinos it follows the total mass distribution however, with less clustering vary the density between these two cases total mass distribution within 100 Mp from peculiar velocity measurements relative overdensities of 2-3

differential proton flux is proportional to the density of relic neutrinos it follows the total mass distribution however, with less clustering vary the density between these two cases total mass distribution within 100 Mp from peculiar velocity measurements relative overdensities of 2-3 printed small sections on A4 papers glued them to a large map read off the density using a ruler

differential proton flux is proportional to the density of relic neutrinos it follows the total mass distribution however, with less clustering vary the density between these two cases total mass distribution within 100 Mp from peculiar velocity measurements relative overdensities of 2-3 printed small sections on A4 papers glued them to a large map read off the density using a ruler

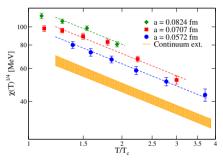






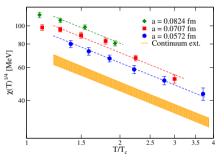
C.Bonati, M.d'Elia, G.Martinelli et al. JHEP 229 03, 155 (2016)

b brute force fully dynamic in the continuum up to $\approx 4 T_c$



C.Bonati, M.d'Elia, G.Martinelli et al. JHEP 229 03, 155 (2016)

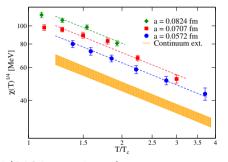
b brute force fully dynamic in the continuum up to $\approx 4 T_c$



Result: $b\sim 3$ unexpected (DIGA etc. $b\sim 8$) for T>2 GeV is larger than DIGA by 7-8 orders of magnitude

C.Bonati, M.d'Elia, G.Martinelli et al. JHEP 229 03, 155 (2016)

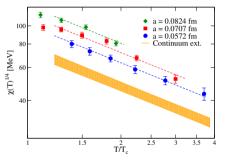
b brute force fully dynamic in the continuum up to $\approx 4 T_c$



Result: $b\sim 3$ unexpected (DIGA etc. $b\sim 8$) for T>2 GeV is larger than DIGA by 7-8 orders of magnitude one order of magnitude shift for the axion dark matter window

C.Bonati, M.d'Elia, G.Martinelli et al. JHEP 229 03, 155 (2016)

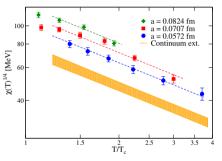
b brute force fully dynamic in the continuum up to $\approx 4T_c$



Result: $b\sim 3$ unexpected (DIGA etc. $b\sim 8$) for T>2 GeV is larger than DIGA by 7-8 orders of magnitude one order of magnitude shift for the axion dark matter window crosses quenched result at $4T_c$ (for quenched $\chi_t^{1/4}(4T_c)$ =17 MeV)

C.Bonati, M.d'Elia, G.Martinelli et al. JHEP 229 03, 155 (2016)

b brute force fully dynamic in the continuum up to $\approx 4T_c$



Result: $b\sim 3$ unexpected (DIGA etc. $b\sim 8$) for T>2 GeV is larger than DIGA by 7-8 orders of magnitude one order of magnitude shift for the axion dark matter window crosses quenched result at $4T_c$ (for quenched $\chi_t^{1/4}(4T_c)$ =17 MeV)

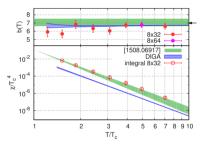
⇒ further study is obviously necessary



Fixed Q simulation: extra acc/rej step at the end of each update, as lattice spacing decreased the acceptance gets better.

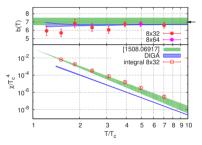
Fixed Q simulation: extra acc/rej step at the end of each update, as lattice spacing decreased the acceptance gets better.

Test in quenched case: pure Wilson action upto $7 \cdot T_c$ and 8×64^3



Fixed Q simulation: extra acc/rej step at the end of each update, as lattice spacing decreased the acceptance gets better.

Test in quenched case: pure Wilson action upto $7 \cdot T_c$ and 8×64^3



standard method: extrapolation using a fit;

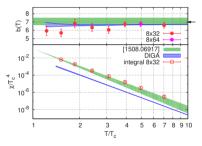
integral method;

Dilute Instanton Gas Approximation:



Fixed Q simulation: extra acc/rej step at the end of each update, as lattice spacing decreased the acceptance gets better.

Test in quenched case: pure Wilson action upto $7 \cdot T_c$ and 8×64^3



standard method: extrapolation using a fit; **integral method**;

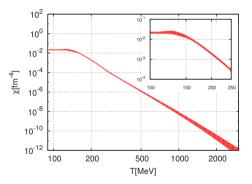
Dilute Instanton Gas Approximation:

exponent agrees nicely, but order of magnitude difference in χ

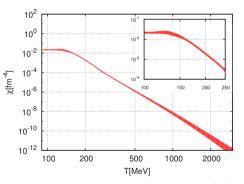


very few topology changes (hard): S. Borsanyi et al. Nature 539 (2016) 69

very few topology changes (hard): S. Borsanyi et al. Nature 539 (2016) 69

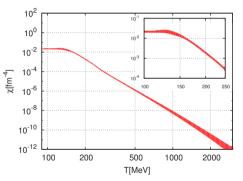


very few topology changes (hard): S. Borsanyi et al. Nature 539 (2016) 69



absolute lower limit (all DM from misalignment): $m_A \gtrsim 28(2)~\mu \mathrm{eV}$

very few topology changes (hard): S. Borsanyi et al. Nature 539 (2016) 69



absolute lower limit (all DM from misalignment): $m_A \gtrsim 28(2)~\mu eV$ assuming 50-99% other (e.g. strings): $m_A = 50 - 1500~\mu eV$